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# Testing asteroseismology with *Gaia* DR2: hierarchical models of the Red Clump

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## ABSTRACT

Asteroseismology provides fundamental stellar parameters independent of distance, but subject to systematics under calibration. *Gaia* DR2 has provided parallaxes for a billion stars, which are offset by a parallax zero-point ( $\varpi_{\text{zp}}$ ). Red Clump (RC) stars have a narrow spread in luminosity, thus functioning as standard candles to calibrate these systematics. This work measures how the magnitude and spread of the RC in the *Kepler* field are affected by changes to temperature and scaling relations for seismology, and changes to the parallax zero-point for *Gaia*. We use a sample of 5576 RC stars classified through asteroseismology. We apply hierarchical Bayesian latent variable models, finding the population-level properties of the RC with seismology, and use those as priors on *Gaia* parallaxes to find  $\varpi_{\text{zp}}$ . We then find the position of the RC, using published values for  $\varpi_{\text{zp}}$ . We find a seismic temperature-insensitive spread of the RC of  $\sim 0.03$  mag in the 2MASS *K* band and a larger and slightly temperature-dependent spread of  $\sim 0.13$  mag in the *Gaia* *G* band. This intrinsic dispersion in the *K* band provides a distance precision of  $\sim 1$  per cent for RC stars. Using *Gaia* data alone, we find a mean zero-point of  $-41 \pm 10 \mu\text{as}$ . This offset yields RC absolute magnitudes of  $-1.634 \pm 0.018$  in *K* and  $0.546 \pm 0.016$  in *G*. Obtaining these same values through seismology would require a global temperature shift of  $\sim -70$  K, which is compatible with known systematics in spectroscopy.

**Key words:** asteroseismology – parallaxes – stars: fundamental parameters – stars: statistics.

## 1 INTRODUCTION

Since the launch of *COROT* (Baglin et al. 2006) and *Kepler* (Borucki, Koch & Team 2010), the use of asteroseismology – the study of stars’ internal physics by observing their modes of oscillation – has become a crucial tool for testing fundamental stellar properties. The large quantity of long time series photometry from these missions (Chaplin & Miglio 2013) and its distance-independent nature have allowed for measures of precise stellar radii and masses for both red giant stars (Hekker et al. 2011; Huber et al. 2011, 2014; Pinsonneault et al. 2014; Mathur et al. 2016; Pinsonneault et al. 2018; Yu et al. 2018) and main sequence

stars (Chaplin et al. 2010, 2011, 2014), studies of exoplanets and exoplanet hosts (Christensen-Dalsgaard et al. 2010; Batalha et al. 2011; Chaplin et al. 2013; Huber et al. 2013a,b; Silva Aguirre et al. 2015), internal and external stellar rotation (Beck et al. 2012; Deheuvels et al. 2012, 2014; Mosser et al. 2012b; Davies et al. 2015), ages of stellar populations (Miglio et al. 2009, 2013; Casagrande et al. 2014, 2016; Stello et al. 2015), and classifications of stellar types (Bedding et al. 2011; Mosser et al. 2012a, 2015; Stello et al. 2013; Vrad, Mosser & Samadi 2016; Elsworth et al. 2017), among others.

Many of these works rely on the so-called ‘direct method’: the use of seismic scaling relations related to the two fundamental oscillation parameters,  $\nu_{\text{max}}$ , the frequency of maximum power of the oscillation mode envelope, and  $\Delta\nu$ , the spacing between two oscillation modes of equal radial degree. These properties are

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individually proportional to mass, radius, and temperature, and, when combined and scaled with solar values, can provide measures of stellar mass, radius, and surface gravity (Kjeldsen & Bedding 1995). As such, stellar properties obtained through seismology depend on temperature as well as on the seismic parameters. Besides the direct method, results from seismology can also be obtained by comparing global seismic properties with a grid of models, referred to as ‘grid modelling’, and can be expanded to ‘detailed modelling’, which directly fits observed seismic mode frequencies to the grids (Metcalfe et al. 2012, 2014; Silva Aguirre et al. 2013, 2015; Davies et al. 2016; Lund et al. 2017).

The seismic scaling relations have been thoroughly tested through interferometry (White et al. 2013), astrometry (Huber et al. 2017), eclipsing binaries (Gaulme et al. 2016), and open clusters (Miglio et al. 2012). Theoretically motivated corrections to the  $\Delta\nu$  and  $\nu_{\max}$  scaling relations have been proposed to depend on  $T_{\text{eff}}$ , metallicity, and evolutionary state (Miglio et al. 2012; Sharma et al. 2016), and it is known that a small correction for the mean molecular weight could be needed for the  $\nu_{\max}$  scaling relation (Belkacem et al. 2013; Viani et al. 2017).

When using the direct method, effective temperatures from spectroscopic analysis are often used (e.g. the APOKASC catalogue; Pinsonneault et al. 2014; Pinsonneault et al. 2018). However, depending on the atmospheric models and temperature scales applied in spectroscopic analysis, inferred values for  $T_{\text{eff}}$  can vary up to  $\sim 170$  K for Core Helium-Burning (CHeB) stars (Slumstrup et al. 2019). While Bellinger et al. (2019) have recently shown that these systematic uncertainties can be mitigated through the use of grid modelling for main-sequence and sub-giant stars, the question of which temperature scale for spectroscopy obtains the best value for  $T_{\text{eff}}$  remains open.

Seismic observations can be combined with distance-dependent observations, such as astrometry, to improve and calibrate results. The second data release (DR2) of the astrometric *Gaia* mission (Gaia Collaboration 2018) recently has provided data for a sample of over one billion targets, with uncertainties largely improved from the first data release (DR1, TGAS; Gaia Collaboration 2016), allowing for a broader range of science and calibrations (Zinn et al. 2018). With DR2, Lindegren et al. (2018) suggested a mean global parallax zero-point offset of  $-29 \mu\text{as}$ , in the sense that *Gaia* parallaxes are too small, using a quasar sample, although it should be noted that the offset varies as a function of colours, magnitude, and position on the sky. Arenou et al. (2018) computed the parallax difference between DR2 and existing catalogues, as well as prior data for individual targets, and found these on average to be the same order of magnitude as the Lindegren et al. (2018) zero-point. Riess et al. (2018) used Cepheid variables to derive a zero-point offset of  $-46 \pm 13 \mu\text{as}$ , Stassun & Torres (2018) used Eclipsing Binaries to find a zero-point of  $-83 \pm 33 \mu\text{as}$ , and Zinn et al. (2018) compared parallaxes to seismic radii to identify a colour- and magnitude-dependent offset of  $-52.8 \pm 2.4(\text{stat.}) \pm 1(\text{syst.}) \mu\text{as}$  for red giant branch (RGB) stars in the *Kepler* field. Finally, using analysis of individual seismic mode frequencies for 93 dwarf stars, Sahlholdt & Silva Aguirre (2018) reported an offset in estimated stellar radii equal to a parallax offset of  $-35 \pm 16 \mu\text{as}$ . As the parallax zero-point offset is known to vary with magnitude, colour, and position in the sky, the differences between these values for the zero-point are expected. Understanding how we quantify the offset is crucial if we want to use *Gaia* to calibrate asteroseismology and other methods.

One method of testing independent sets of measurements is calculating an observable astronomical property. An example of

such a property is the luminosity of the ‘Red Clump’ (RC), an over-density of red giant stars on the HR-diagram, in bands of absolute magnitude. When stars of masses around  $0.7 \lesssim M/M_{\odot} \lesssim 1.9$  (for  $[M/H] \simeq 0.07$ , upper limit subject to change with metallicity) ignite helium in their cores, they undergo the He-flash. The He-burning core masses are very similar for these stars, and as their luminosity is mainly determined by the core mass, they will all have similar luminosities, creating a clump of stars on the HR-diagram (see Girardi 2016, and references therein). Further differences in luminosity and temperature are then effects of metallicity and envelope mass, and thus the Clump has a relatively small spread. Stars at lower masses and low metallicities form a horizontal branch at a luminosity similar to the RC, whereas stars of masses just above the limit for the He-flash lie at a slightly lower luminosity, forming a Secondary Red Clump (2CL; Girardi 1999). At even higher masses, the luminosity becomes a function of stellar mass, and these stars form a vertical structure in the HR-diagram during their CHeB phase.

The luminosity of the RC overdensity may be used as a standard candle, given constraints on mass and metallicity (Cannon 1970), and has recently been used to calibrate *Gaia* DR1 parallaxes (Davies et al. 2017). Also using *Gaia* DR1 parallaxes, Hawkins et al. (2017) (hereafter H17) found precise measurements for the RC luminosity in various passbands, including the 2MASS *K* band, which minimized the spread in luminosity due to mass and metallicity (Salaris & Girardi 2002). With *Gaia* DR2’s improved parallax uncertainties and reduced systematic offset, now is a good time to revisit the RC as a calibrator.

In this work we investigate systematics in both asteroseismology and *Gaia* simultaneously, to see how differences in assumptions for one influence inferences of the other. Using a sample of over 5500 *Kepler* Red Giant stars in the RC for which parallaxes and seismology are available, we measure the position of the RC population in absolute magnitude in the 2MASS (Skrutskie et al. 2006) *K* band and the *Gaia* *G* band. We do this using seismology and parallax (with photometry) independently. Since the distribution of RC stars should be the same for this population, independent of method, a (dis)agreement of the measured positions and spreads of the RC using two independent methods sheds light on systematics in both. For the seismic method, we test the influence of the temperature scale used to obtain the values of  $T_{\text{eff}}$  fed into seismic scaling relations, as well as the impact of corrections to the  $\Delta\nu$  scaling relation. For the *Gaia* method, we study how changes in the parallax zero-point offset for *Gaia* DR2 impact the inferred luminosity of the RC.

This paper is laid out as follows. Section 2 discusses how the data were obtained, and the theory used to calculate our observables. Section 3 discusses how we use hierarchical Bayesian modelling to study the RC. We present our results in Section 4 and discuss them in context of similar work in Section 5, and present our conclusions in Section 6.

## 2 DATA

Our aim is to find the intrinsic position and spread of the RC in absolute magnitude for various passbands using two approaches: one using a distance-independent luminosity calculated from asteroseismology and the other using a magnitude inferred from photometry and *Gaia* DR2 parallaxes. Since the number of stars with asteroseismic data is significantly lower than those with data in *Gaia* DR2, this limits our sample.

For our asteroseismic sample, we used the catalogue of 16 094 oscillating *Kepler* red giants by Yu et al. (2018) (hereafter Y18),

which contains global oscillation parameters  $\nu_{\max}$  and  $\Delta\nu$ , as well as broad evolutionary state classifications, effective temperatures  $T_{\text{eff}}$ , and metallicities  $[\text{Fe}/\text{H}]$  taken from Mathur et al. (2017).

We re-considered the classification of all stars labelled as CHeB in the Y18 catalogue, using the method presented in Elsworth et al. (2017). This uses the structure of dipole-mode oscillations in the power spectra to classify stars as belonging to the 2CL, the RGB, or the RC. We obtained light curves for 7437 stars labelled as CHeB in Y18, from two sources: the so-called KASOC light curves (Handberg & Lund 2014)<sup>1</sup> and the KEPSEISMIC light curves (García et al. 2011).<sup>2</sup> The latter have been produced with larger photometric masks to ensure a better stability at low frequencies, and have been gap-filled using in-painting techniques (García et al. 2014; Jofré et al. 2015).

Of these 7437 stars, we found that 5668 are RC, 737 are 2CL, and no classification could be found for 499 stars. Notably, 533 stars were found to be RGB, disagreeing with the classification listed in Y18. This should be discussed in future work, but for the sake of internal consistency of our classifications, we have chosen to adopt the Elsworth et al. (2017) classification in this work.

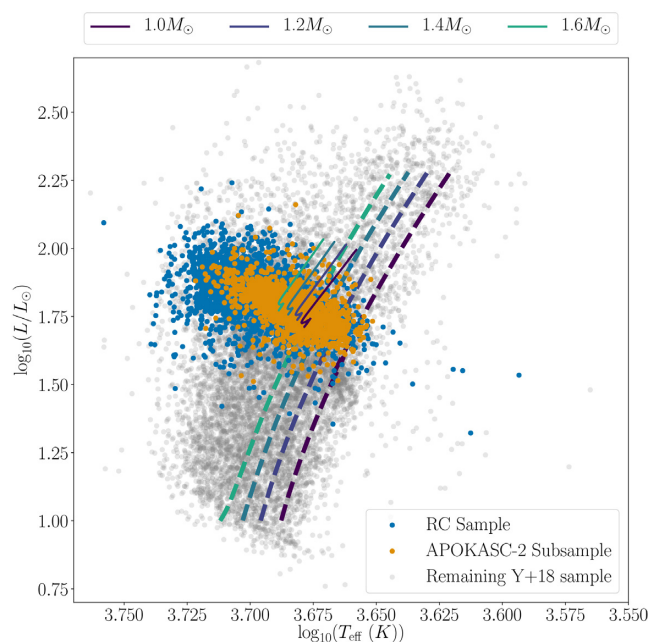
It should be noted that our classification does not specifically account for low-mass, low-metallicity horizontal branch stars, which are therefore expected to be retained in our sample, but are not expected to significantly affect the result as they have similar luminosities to the RC, and no extensive horizontal structure is present on the HR diagram of the Y18 catalogue, or our subsample thereof (see Fig. 1). A fraction of the newly classified stars had masses reported in Y18 as much higher than we would expect for an RC star. In order to exclude these from our sample, we apply a liberal cut for clump-corrected seismic masses of over  $2.2 M_{\odot}$ , excluding 92 stars from our sample.

To obtain our astrometric sample, we cross-matched the RC stars we selected from the Y18 sample with the *Gaia* DR2 sample<sup>3</sup> (Gaia Collaboration 2016, 2018). In cases of duplicate sources for a given KIC, we selected the star with the lowest angular separation to the target. We did not apply any truncation of the sample based on parallax uncertainty or negative parallax, since this is known to introduce a parallax-dependent bias (Luri et al. 2018).

The parallaxes ( $\hat{\pi}$ ) and parallax uncertainties ( $\sigma_{\hat{\pi}}$ ) make up our astrometric set of observables. We obtained the apparent magnitudes ( $\hat{m}$ ) and their uncertainties ( $\sigma_{\hat{m}}$ ) from the 2MASS survey for the *K* band (Skrutskie et al. 2006) and *Gaia* DR2 for the *Gaia* *G* band, and removed stars that do not have photometry or uncertainties on magnitude in 2MASS. Comparing the magnitude zero-points for the *Gaia* *G*, *G<sub>BP</sub>* and *G<sub>RP</sub>* bands, Casagrande & VandenBerg (2018b) found indication of a magnitude-dependent zero-point offset in the *Gaia* *G*-band magnitudes in the range of  $6 \text{ mag} \lesssim G \lesssim 16.5 \text{ mag}$ , corrected as

$$G^{\text{corr}} = 0.0505 + 0.9966 G, \quad (1)$$

where  $G$  is our uncorrected *Gaia* *G*-band magnitude. This correction is small, and corresponds to 30 mmag over 10 magnitudes. We gave all our *G*-band magnitudes a generous uncertainty of 10 mmag, the typical uncertainty quoted in Gaia Collaboration (2018) for  $G = 20$ ,



**Figure 1.** HR diagram illustrating the data in our final set of 5576 stars overlaid on the Y18 sample, along with evolutionary tracks from MESA (Paxton et al. 2011, 2013, 2015) (for details about the physical inputs of the models, see Khan et al. 2018). The stars in the Y18 sample not in our final selection are in grey. Plotted on top in blue are the stars that in our final sample where the subsample of stars with temperatures reported in APOKASC-2 (Pinsonneault et al. 2018) are shown in orange. Evolutionary tracks are plotted for masses ranging between 1.0 and 1.6 solar masses for a metallicity of  $Z = 0.01108$  and helium content of  $Y = 0.25971$ . The dashed lines indicate the RGB, whereas the solid lines indicate the main Core Helium Burning stage of the tracks [the Helium flash (and subflashes) are not included].

in order to account for any additional uncertainty incurred by the above correction. It should be noted that a similar relation for the correction of *G*-band magnitudes is presented in Maíz Apellániz & Weiler (2018). This correction places magnitudes about 30 mmag higher than when using the Casagrande & VandenBerg (2018b) correction in the applicable magnitude range. We expect the scale of this systematic offset to have a negligible impact on our results, and therefore adopt the Casagrande & VandenBerg (2018b) correction in this work for consistency with our chosen *G*-band extinction coefficients and bolometric corrections (BCs).

Our model also uses an extinction for each star in each band. Reddening values are taken from the Green et al. (2018) three-dimensional dustmap under the assumption that the distance to the object is that given by Bailer-Jones et al. (2018). We note that this is not expected to bias our results towards a previous measure of distance, because the spread in the obtained reddening values, regardless of choice of distance value, falls well within the spread of the prior set on these values in our model. We converted reddening to the band-specific extinction  $\hat{A}_{\lambda}$  using extinction coefficients unique to the Green et al. (2018) map for the *K* band.<sup>4</sup> For the *Gaia* *G* band we calculated our band-specific extinction using the mean extinction coefficient presented in Casagrande & VandenBerg (2018b), after converting our reddening value to a measure of  $E(B - V)$  following the conventions presented in Green et al. (2018).

<sup>4</sup>These coefficients can be found with the Green et al. (2018) usage notes.

<sup>1</sup>Freely distributed at the KASOC webpage (<http://kasoc.phys.au.dk>).

<sup>2</sup>Freely distributed at the MAST website (<https://archive.stsci.edu/prepds/kepseismic/>).

<sup>3</sup>We make use of the of the <https://www.gaia-kepler.fun> crossmatch database created by Megan Bedell for this purpose.



The final sample contains 5576 RC stars, with minimal contamination from the 2CL or the RGB, and covers a magnitude range of  $\sim 8$  to  $\sim 16$  mag in  $G$  and  $\sim 6$  to  $\sim 14$  mag in  $K$ . Note that for this magnitude range we expect the *Gaia* DR2 catalogue to be practically complete, and do not need to apply any selection functions in magnitude. The data are shown in Fig. 1 in an HR diagram overlaid on the full Y18 sample.

## 2.1 The APOKASC-2 subsample

We used temperatures from Mathur et al. (2017), a catalogue compiling temperatures from a diverse set of papers including work with spectroscopy, photometry, and some asteroseismology. In order to investigate the impact of using differing temperature sources on our results, we also included runs on a subsample of 1637 stars that had  $T_{\text{eff}}$  values reported in the APOKASC-2 catalogue (Pinsonneault et al. 2014; Pinsonneault et al. 2018). When calculating seismic properties from these data, we only changed the values for  $T_{\text{eff}}$  to our new APOKASC-2 values. In Fig. 2 we compare the distributions in  $T_{\text{eff}}$ , mass, radius, and  $[\text{Fe}/\text{H}]$  of the Y18 RC sample and the APOKASC-2 subsample. Also shown is the distribution of the APOKASC-2 temperatures, which are overall lower than the Y18 temperatures, and the distributions in mass and radius calculated through the direct method for these temperatures. Overall the APOKASC-2 subsample represents a lower temperature population, with its most distinct difference being in  $T_{\text{eff}}$  and  $[\text{Fe}/\text{H}]$ .

## 2.2 Obtaining the seismic sample

The two global observable seismic parameters,  $\nu_{\text{max}}$  and  $\Delta\nu$ , scale with fundamental stellar properties as (Brown et al. 1991; Kjeldsen & Bedding 1995):

$$\frac{\nu_{\text{max}}}{\nu_{\text{max}\odot}} \simeq \left(\frac{M}{M_{\odot}}\right) \left(\frac{R}{R_{\odot}}\right)^{-2} \left(\frac{T_{\text{eff}}}{T_{\text{eff}\odot}}\right)^{-1/2} \quad \text{and} \quad (2)$$

$$\frac{\Delta\nu}{\Delta\nu_{\odot}} \simeq \left(\frac{M}{M_{\odot}}\right)^{1/2} \left(\frac{R}{R_{\odot}}\right)^{-3/2}, \quad (3)$$

where  $M$  is the stellar mass,  $R$  is the radius,  $T_{\text{eff}}$  is the effective temperature, and  $\odot$  indicates a solar value. In this work we used  $\nu_{\text{max}\odot} = 3090 \pm 30 \mu\text{Hz}$ ,  $\Delta\nu_{\odot} = 135.1 \pm 0.1 \mu\text{Hz}$ , and  $T_{\text{eff}\odot} = 5777 \text{ K}$  (Huber et al. 2011). By rearranging these scaling relations, we can obtain stellar surface gravity and radius as

$$\frac{g}{g_{\odot}} \simeq \frac{\nu_{\text{max}}}{\nu_{\text{max}\odot}} \left(\frac{T_{\text{eff}}}{T_{\text{eff}\odot}}\right)^{1/2} \quad \text{and} \quad (4)$$

$$\frac{R}{R_{\odot}} \simeq \left(\frac{\nu_{\text{max}}}{\nu_{\text{max}\odot}}\right) \left(\frac{\Delta\nu}{f_{\Delta\nu} \Delta\nu_{\odot}}\right)^{-2} \left(\frac{T_{\text{eff}}}{T_{\text{eff}\odot}}\right)^{1/2}, \quad (5)$$

where the new term  $f_{\Delta\nu}$  is a correction to the  $\Delta\nu$  scaling relation in the notation of Sharma et al. (2016). We calculated  $f_{\Delta\nu}$  as a function of  $[\text{Fe}/\text{H}]$ ,  $T_{\text{eff}}$ ,  $\nu_{\text{max}}$ ,  $\Delta\nu$  and evolutionary state using interpolation in a grid of models (Sharma & Stello 2016). For each perturbation of  $T_{\text{eff}}$ , we recalculated  $f_{\Delta\nu}$ , changing no other parameters. We only extracted the correction values  $f_{\Delta\nu}$  from the models, and used the seismic parameters and temperature values from our original set, and not the results for these values returned from the grids, in the rest of this work. We did not include corrections for the  $\nu_{\text{max}}$  scaling relation, because these are more difficult to obtain theoretically (Belkacem et al. 2011), and are probably negligible (Brogaard et al. 2018). Note that Brogaard et al. (2018) found that using corrections

by Rodrigues et al. (2017) delivers on average slightly smaller stellar properties than using Sharma & Stello (2016) due to differences in how these methods treat the solar surface effect. Since we used a wide range of BCs for various temperature perturbations, the method by Rodrigues et al. (2017) would be too computationally expensive, and we thus elected to use Sharma & Stello (2016), which may lead to differences of the order of  $\sim 2$  per cent in radius than if we had used Rodrigues et al. (2017) (White et al. 2011). We discuss the impact of this on our work in Section 5.

In order to obtain absolute magnitudes for our sample, we used  $T_{\text{eff}}$  and seismic radii, calculated using  $\Delta\nu$  and  $\nu_{\text{max}}$  from the Y18 catalogue through equation (5), to calculate the stellar luminosity as

$$L_* = 4\pi\sigma_{\text{sb}} R^2 T_{\text{eff}}^4. \quad (6)$$

Here  $L_*$  is the luminosity of the star and  $\sigma_{\text{sb}}$  is the Stefan–Boltzmann constant. This was converted to a bolometric magnitude as in Casagrande & VandenBerg (2014):

$$M_{\text{bol}} = -2.5 \log_{10}(L_*/L_{\odot}) + M_{\text{bol}\odot}, \quad (7)$$

where  $L_{\odot}$  is the solar luminosity, and we have adopted  $M_{\text{bol}\odot} = 4.75$  (Casagrande & VandenBerg 2014, 2018a,b). We calculated the BC in the 2MASS  $K$  and *Gaia*  $G$  bands with the method described by Casagrande & VandenBerg (2014, 2018a,b) using  $T_{\text{eff}}$ ,  $[\text{Fe}/\text{H}]$  and  $\log g$ , and without accounting for extinction. Since we are using a distance-independent measure of luminosity to calculate an absolute magnitude, accounting for this in the BC would bias our results. Because our method requires tweaking our values for  $T_{\text{eff}}$ , we recalculated the  $\log g$  used to find the BC through the scaling relation in equation (4), as well as our values for  $f_{\Delta\nu}$ , for each different set of temperatures, and thus obtained a full set of BCs and corrections to the scaling relations for each temperature perturbation. Our values of absolute magnitude were then given by

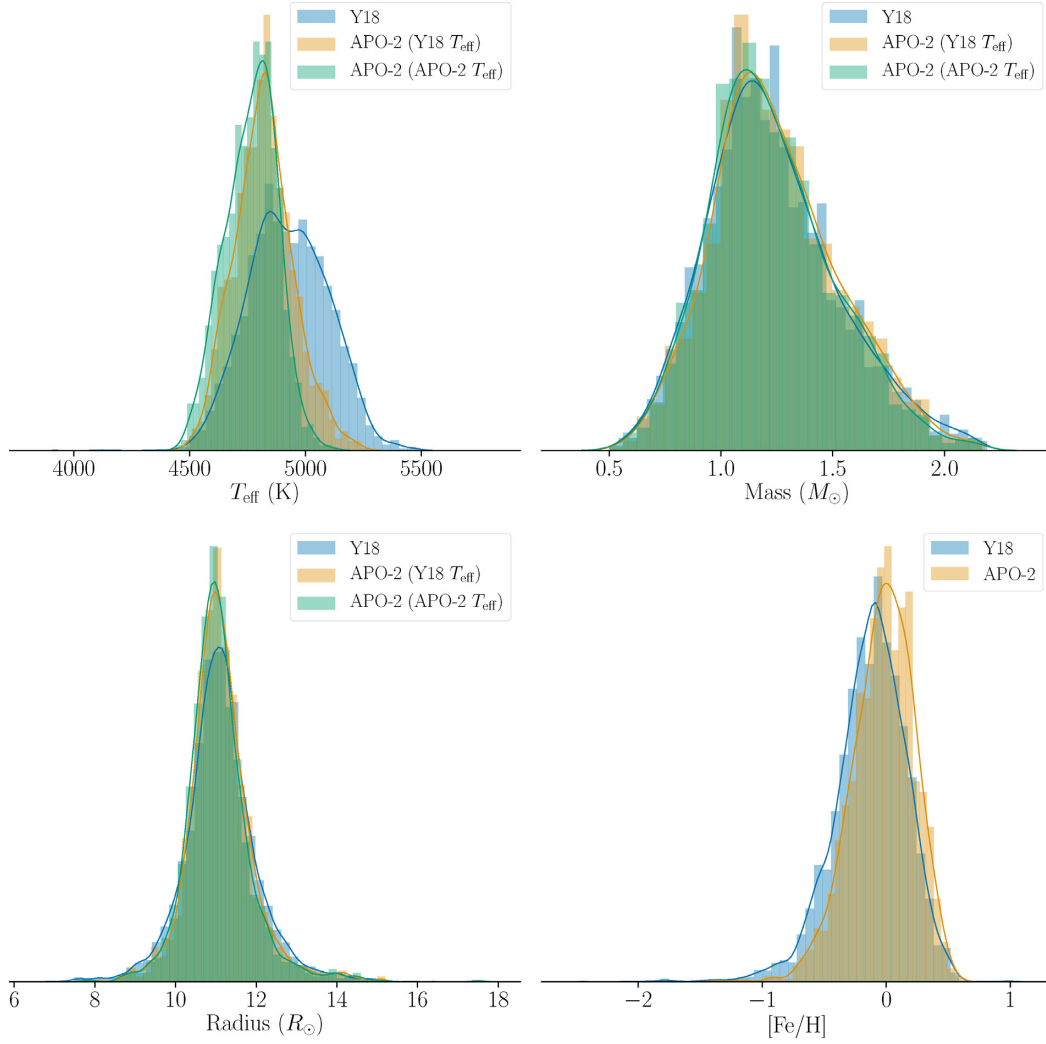
$$\hat{M}_{\lambda} = M_{\text{bol}} - \text{BC}_{\lambda}, \quad (8)$$

where  $\lambda$  is the relevant band,  $M_{\text{bol}}$  is the bolometric luminosity, and  $\text{BC}_{\lambda}$  is the bolometric correction in that band. Uncertainties on  $\hat{M}_{\lambda}$  were propagated through from the uncertainties on seismic parameters and effective temperatures, including those on the solar seismic parameters. Uncertainties on the BCs were estimated using a Monte Carlo method with 5000 iterations for 1000 randomly selected stars from our sample. We found an uncertainty of 0.3 mag for all BCs in the  $G$  band. For the  $K$  band we found 0.05 mag for stars with a fractional temperature uncertainty of  $< 2.5$  per cent, and 0.09 mag for those with larger fractional uncertainties on temperature. We discuss the systematic uncertainties on  $f_{\Delta\nu}$  in Section 5.

## 3 LOCATING THE RED CLUMP USING HIERARCHICAL BAYESIAN MODELLING

In order to test systematics in asteroseismology and *Gaia* using the RC, we aim to find the location and spread of the RC in absolute magnitude using both sets of data separately. To obtain these RC parameters, we fitted a model for the distribution of RC stars in ‘true’ absolute magnitude, either inferred from an observed absolute magnitude (asteroseismic) or inferred from apparent magnitude, parallax, and extinction (astrometric).

We built a pair of Bayesian hierarchical models with latent parameters that allow us to infer key values such as the distance and the true absolute magnitude from the data and the model. The latent parameters form a stepping stone between our population model, which is described by *hyperparameters*, and the observations. We use a latent parameter for each star to infer the ‘true’ distribution of



**Figure 2.** Distributions in  $T_{\text{eff}}$ , mass, radius, and  $[\text{Fe}/\text{H}]$  of the RC sample (Yu et al. 2018) and the APOKASC-2 subsample (Pinsonneault et al. 2014, 2018). In green are the distribution of the APOKASC-2 temperatures, which are overall lower, and the distributions in mass and radius calculated through the direct method for these temperatures. In the labels, ‘APO-2’ represents a shorthand for APOKASC-2.

the absolute magnitudes, while fitting our population level model to these inferred ‘true’ absolute magnitudes, instead of to the observations themselves. Many aspects of our hierarchical models, especially those for the *Gaia* data, are similar to those used for the same purpose by H17 with some improvements.

To fit to the position and spread of RC stars while also isolating any outlier contaminants, we applied the mixture model (Hogg, Bovy & Lang 2010) utilized by H17. In this case, we employed two generative models weighted by the mixture-model weighting factor  $Q$ . For these we used two normal distributions: one for the inlier population of RC stars, with a mean  $\mu_{\text{RC}}$  and a standard deviation (spread)  $\sigma_{\text{RC}}$ , and a broad outlier distribution centred in the same location ( $\mu_{\text{RC}}$ ) but with a spread of  $\sigma_o$ , which must always be larger than  $\sigma_{\text{RC}}$ . The likelihood to obtain an absolute magnitude  $M_i$  given this mixture model is then

$$p(M_i|\theta_{\text{RC}}) = Q\mathcal{N}(M_i|\mu_{\text{RC}}, \sigma_{\text{RC}}) + (1 - Q)\mathcal{N}(M_i|\mu_{\text{RC}}, \sigma_o), \quad (9)$$

where  $M_i$  is the true absolute magnitude for a given datum  $i$ ,  $\theta_{\text{RC}} = \{\mu_{\text{RC}}, \sigma_{\text{RC}}, Q, \sigma_o\}$  are the model hyperparameters (which inform

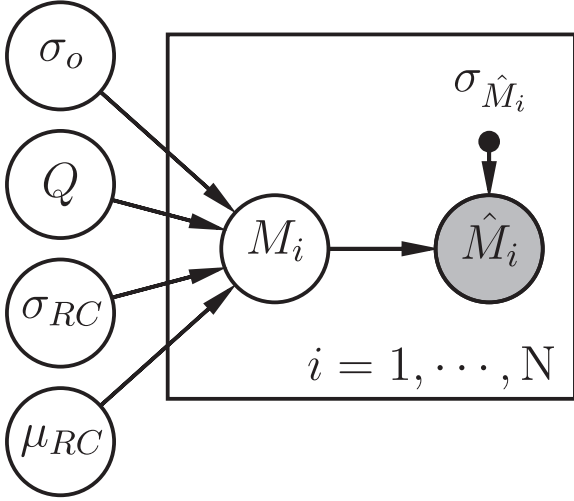
the population of latent parameters) and  $\mathcal{N}(x|\mu, \sigma)$  represents a normal distribution evaluated at  $x$ , with a mean  $\mu$  and a spread  $\sigma$ .<sup>5</sup>

### 3.1 The asteroseismic model

For our asteroseismic model, we used a calculated measure of the absolute magnitude ( $\hat{M}$ ) from asteroseismology, along with appropriate uncertainties ( $\sigma_{\hat{M}}$ ), as our data. We used a latent variable model to infer the true value of the absolute magnitude. Given our data and the hyperparameters on our mixture model  $\theta_{\text{RC}}$ , we can use Bayes’ theorem to find the unnormalized posterior probability of our model:

$$p(\theta_{\text{RC}}|\mathcal{D}) \propto p(\theta_{\text{RC}}) \prod_{i=1}^N p(\mathcal{D}_i|M_i)p(M_i|\theta_{\text{RC}}). \quad (10)$$

<sup>5</sup>Note that the spread  $\sigma$  as listed in  $\mathcal{N}(x|\mu, \sigma)$  is not a variance, but a standard deviation, since we are following the nomenclature used in *pystan*.



**Figure 3.** A probabilistic graphical model of the astroseismic model, represented algebraically in equation (10). Shaded circles indicate observed data, whereas solid black circles represent fixed parameters, such as the uncertainty on the observed data. The hyperparameters  $\theta_{RC}$  can be seen on the left, and inform the set of latent parameters  $M_i$ , which in turn relate to the observed data  $\hat{M}_i$  and  $\sigma_{\hat{M}_i}$ .  $N$  is the number of data points in our sample.

Here,  $N$  is the number of points in our data set  $\mathcal{D} = \{\hat{M}, \sigma_{\hat{M}}\}$ ,  $p(\mathcal{D}_i | M_i)$  is our likelihood function,  $p(\theta_{RC})$  represents the priors on the hyperparameters, and  $p(M_i | \theta_{RC})$  is the probability to obtain our latent parameters (the true absolute magnitudes) given our hyperparameters.

The likelihood to obtain our data given our parameters is then

$$p(\mathcal{D}_i | M_i) = \mathcal{N}(\hat{M}_i | M_i, \sigma_{\hat{M}_i}), \quad (11)$$

where  $M_i$  is the true absolute magnitude. Here,  $M_i$  is a latent parameter that is drawn from the likelihood function  $p(M_i | \theta_{RC})$  (equation 9), to which our hyperparameters are fitted. A probabilistic graphical model of the astroseismic model is shown in Fig. 3.

### 3.2 The astrometric model

Fitting the absolute magnitude for the *Gaia* DR2 sample required a more involved approach, since we wanted to work directly with parallax (Luri et al. 2018). We used a set of three latent parameters,  $\alpha_i = \{M_i, r_i, A_i\}$ , where  $M_i$  is the absolute magnitude in a given band,  $r_i$  is the distance, and  $A_i$  is the extinction in a given band. We also include two additional hyperparameters:  $\varpi_{zp}$ , the parallax zero-point offset, and  $L$ , the length scale of the exponentially decreasing space density prior on distance (Astraatmadja & Bailer-Jones 2016a,b, 2017). This prior, which is necessary to treat negative parallax values, has already successfully been applied to *Gaia* DR2 data (Bailer-Jones et al. 2018) and its use is recommended for this purpose within the *Gaia* DR2 release papers (Luri et al. 2018).

Some extra care was also required in the treatment of parallax uncertainties for this sample. Lindegren et al. (2018) found parallaxes to be correlated on scales below  $40^\circ$ , with increasing strength at smaller separations, and quantified their covariance using quasar parallaxes. They found the positive covariance  $V_{\varpi}$  for these scales to be reasonably approximated by the fitted relation

$$V_{\varpi}(\theta) \simeq (285 \mu\text{as}^2) \times \exp(-\theta/14^\circ), \quad (12)$$

where  $\theta$  is the angular separation between two targets in degrees. The fit corresponds to an RMS amplitude of  $\sqrt{285 \mu\text{as}^2} \approx 17 \mu\text{as}$ .

This relation was recently applied by Zinn et al. (2018), who found that the Lindegren et al. (2018) relation resulted in the best goodness of fit for their models of the parallax zero-point offset, over both a similar relation by Zinn et al. (2017) based on TGAS data, and not including parallax covariances altogether.

We generated a covariance matrix  $\underline{\Sigma}$  for our sample:

$$\Sigma_{ij} = V_{\varpi}(\theta_{ij}) + \delta_{ij} \sigma_{\hat{\varpi}_i} \sigma_{\hat{\varpi}_j}, \quad (13)$$

where  $\theta_{ij}$  is the angular separation between stars  $i$  and  $j$ , and  $\delta_{ij}$  is the Kronecker delta function.

Given these new additions, our set of data was  $\mathcal{D} = \{\hat{\varpi}, \underline{\Sigma}, \hat{m}, \sigma_{\hat{m}}, \hat{A}\}$ , where all symbols are as defined above and  $\hat{A}$  is the band-specific extinction. We can use Bayes' theorem, as before, to find the unnormalized posterior probability of our model as

$$p(\theta_{RC}, \varpi_{zp}, L, \alpha | \mathcal{D}) \propto p(\theta_{RC}, \varpi_{zp}, L, \alpha) p(\mathcal{D} | \theta_{RC}, \varpi_{zp}, L, \alpha), \quad (14)$$

where  $p(\mathcal{D} | \theta_{RC}, \varpi_{zp}, L, \alpha)$  is now our likelihood function and  $p(\theta_{RC}, \varpi_{zp}, L, \alpha)$  represents the priors on our hyper- and latent parameters. Our likelihood function relates to two observables as

$$p(\mathcal{D} | \theta_{RC}, \varpi_{zp}, L, \alpha) = p(\hat{\varpi} | r, \varpi_{zp}, \underline{\Sigma}) \times p(\hat{m} | \alpha, \sigma_{\hat{m}}). \quad (15)$$

Note that the parallax only depends on the latent parameter for distance,  $r$ . Since parallax values are correlated,  $p(\hat{\varpi} | r, \varpi_{zp}, \underline{\Sigma})$  was evaluated for all data simultaneously, whereas  $p(\hat{m} | \alpha, \sigma_{\hat{m}})$  was evaluated at every datum  $i$ . This means that our full posterior probability takes the form

$$p(\theta_{RC}, \varpi_{zp}, L, \alpha | \mathcal{D}) \propto p(\theta_{RC}, \varpi_{zp}, L) p(\hat{\varpi} | r, \varpi_{zp}, \underline{\Sigma}) \times \prod_{i=1}^N p(\hat{m}_i | \alpha_i, \sigma_{\hat{m}_i}) p(\alpha_i | \theta_{RC}, \varpi_{zp}, L), \quad (16)$$

where the first term represents the priors on our hyperparameters, the second term is the likelihood to obtain our observed parallaxes, the third is the likelihood to obtain an observed magnitude, and the fourth gives the probability to obtain the latent parameters, given the hyperparameters.

The second component of equation (16) is the probability of obtaining the observed parallax given our latent parameters and our covariance matrix. Since we treated our parallax uncertainties as correlated, we evaluated these probabilities for the full set using a multivariate normal distribution:

$$p(\hat{\varpi} | r, \varpi_{zp}, \underline{\Sigma}) = \mathcal{N}(\hat{\varpi} | 1/r + \varpi_{zp}, \underline{\Sigma}), \quad (17)$$

where  $1/r$  defines the *true* parallax. The latent parameters for the distance  $r_i$  were drawn from an exponentially decreasing space density prior (Bailer-Jones 2015), which goes as

$$p(r_i | L) = \frac{1}{2L^3} r_i^2 \exp(-r_i/L), \quad (18)$$

and thus depends on the length-scale hyperparameter  $L$ . This prior has a mode at  $2L$ , beyond which it decreases exponentially.

The third component of equation (16) is then

$$p(\hat{m}_i | \alpha_i, \sigma_{\hat{m}_i}) = \mathcal{N}(\hat{m}_i | m_i, \sigma_{\hat{m}_i}), \quad (19)$$

where  $m_i$  is the *true* apparent magnitude, and is drawn from the relation

$$m_i = M_i + 5 \log_{10}(r_i) - 5 + A_i. \quad (20)$$





**Table 1.** Medians of the posterior distributions for hyperparameters of our seismic model, for the 2MASS *K* band, for 5576 stars from the Y18 sample. Uncertainties are taken as the  $1\sigma$  credible intervals, and are listed as a single value for cases where the posterior was approximately Gaussian. Values are listed for data that have been left uncorrected (no correction) and data with an appropriate correction to the seismic scaling relations (Clump Corrected).  $\Delta T_{\text{eff}}$  is the global shift to our values of  $T_{\text{eff}}$ ,  $\mu_{\text{RC}}$  is the position of the RC in absolute magnitude,  $\sigma_{\text{RC}}$  is the spread of the RC in absolute magnitude,  $Q$  is the mixture model weighting factor (and the effective fraction of stars considered inliers), and  $\sigma_o$  is the spread of our outlier population, expressed in terms of  $\sigma_{\text{RC}}$ .

$\Delta T_{\text{eff}}$ (K)	No correction				Clump corrected			
	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )
−50.0	−1.704 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.35 <sup>+1.17</sup> <sub>−1.01</sub>	−1.713 ± 0.002	0.034 ± 0.004	0.91 ± 0.01	8.85 <sup>+1.09</sup> <sub>−0.93</sub>
−40.0	−1.709 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.33 <sup>+1.22</sup> <sub>−1.01</sub>	−1.718 ± 0.002	0.033 ± 0.004	0.91 ± 0.01	9.11 <sup>+1.12</sup> <sub>−1.04</sub>
−30.0	−1.714 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.4 <sup>+1.15</sup> <sub>−1.04</sub>	−1.724 ± 0.002	0.033 ± 0.004	0.91 ± 0.01	9.16 <sup>+1.12</sup> <sub>−0.96</sub>
−20.0	−1.719 ± 0.002	0.029 ± 0.003	0.92 ± 0.01	10.55 <sup>+1.15</sup> <sub>−1.05</sub>	−1.73 ± 0.002	0.033 ± 0.004	0.91 ± 0.01	9.22 <sup>+1.05</sup> <sub>−0.91</sub>
−10.0	−1.724 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.49 <sup>+1.13</sup> <sub>−1.03</sub>	−1.735 ± 0.002	0.033 ± 0.004	0.91 ± 0.01	9.16 <sup>+1.09</sup> <sub>−0.98</sub>
0.0	−1.729 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.33 <sup>+1.19</sup> <sub>−1.01</sub>	−1.741 ± 0.002	0.033 ± 0.004	0.91 ± 0.01	9.18 <sup>+1.09</sup> <sub>−0.92</sub>
10.0	−1.734 ± 0.002	0.029 ± 0.003	0.92 ± 0.01	10.44 <sup>+1.07</sup> <sub>−0.97</sub>	−1.746 ± 0.002	0.032 ± 0.004	0.91 ± 0.01	9.36 <sup>+1.2</sup> <sub>−1.05</sub>
20.0	−1.739 ± 0.002	0.03 ± 0.004	0.92 ± 0.01	10.32 <sup>+1.17</sup> <sub>−1.02</sub>	−1.752 ± 0.002	0.033 ± 0.004	0.91 ± 0.01	9.19 <sup>+1.16</sup> <sub>−1.01</sub>
30.0	−1.744 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.41 <sup>+1.06</sup> <sub>−0.99</sub>	−1.757 ± 0.002	0.032 ± 0.004	0.91 ± 0.01	9.37 <sup>+1.16</sup> <sub>−1.02</sub>
40.0	−1.749 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.41 <sup>+1.18</sup> <sub>−1.02</sub>	−1.762 ± 0.002	0.032 ± 0.004	0.91 ± 0.01	9.37 <sup>+1.14</sup> <sub>−0.97</sub>
50.0	−1.754 ± 0.002	0.03 ± 0.003	0.92 ± 0.01	10.27 <sup>+1.12</sup> <sub>−1.01</sub>	−1.768 ± 0.002	0.032 ± 0.004	0.91 ± 0.01	9.25 <sup>+1.16</sup> <sub>−1</sub>

approximately Gaussian, we quote a symmetric single uncertainty. The change of the posterior on the magnitude of the RC  $\mu_{\text{RC}}$  alone, given the input, can be seen in Fig. 5.

For our APOKASC-2 temperature subsample of 1637 stars, we reran our models using the same methodology as before, simply substituting the temperatures and temperature uncertainties reported in Pinsonneault et al. (2018) for those in Y18 for those stars, and making no other changes. Note that the change in temperature values carried through to the calculation of the BCs and corrections to the scaling relations for each run. The results of this are presented in Tables 2 and 4 for all hyperparameters, as with the run on the full sample. The change in the posteriors on the position of the RC is shown for this reduced sample in Fig. 6.

## 4.2 Results from Gaia

Given our results from asteroseismology, we wish to determine the parallax zero-point offset,  $\varpi_{\text{zp}}$ , that recovers our values of the absolute magnitude and spread of the RC. Since  $\mu_{\text{RC}}$  and  $\sigma_{\text{RC}}$  represent astrophysical observables that should be consistent across both data sets, we used a description of the posterior distributions from these parameters from our seismic model as a highly informative prior in our *Gaia* model. This yields the parallax offset required to recover the same magnitude and spread of the RC found using seismology. We passed in the seismic posteriors for  $\Delta T_{\text{eff}}$  being −50, 0, and +50 K from our runs on our full sample and the reduced APOKASC-2 sample, and thus ran our model for six different RC magnitudes and spreads in each band. Additionally, we used the median values of each latent parameter  $M_i$  from the application of our seismic model to our full sample, along with distance estimates by Bailer-Jones et al. (2018) and observed extinctions from Green et al. (2018), as initial guesses in our *Gaia* model for computational efficiency. No other values were changed on each run.

Following the relation presented in equation (12) (Lindgren et al. 2018), we treated our parallax uncertainties as correlated as a function of position on the sky across the entire *Kepler* field, similarly to previous work by Zinn et al. (2018). While the model equation presented by Lindgren et al. (2018) describes the covariance well for a wide range of separations, individual covariances oscillate

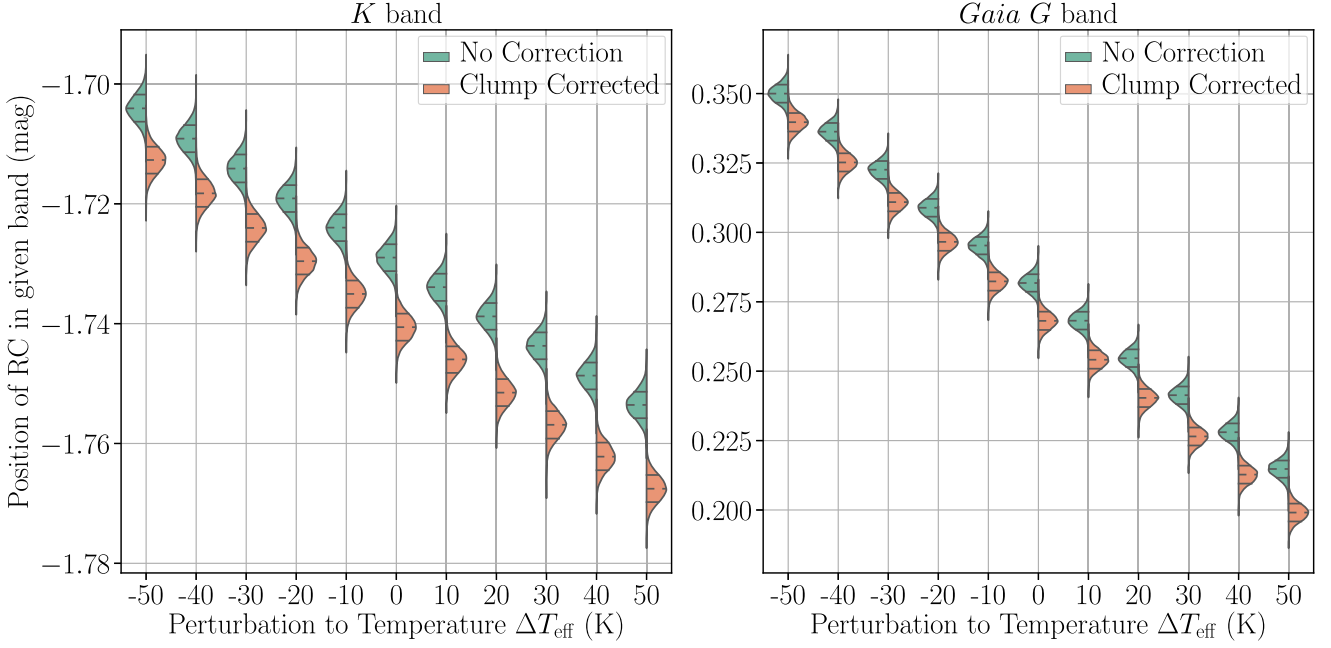
around the model at separations below 1 deg, and the model no longer holds at all for separations below 0.125 deg. To ensure that our treatment of the parallax covariances was sensible, we ran our *Gaia* model on a reduced sample of 1000 stars, randomly selected from across the entire *Kepler* field to ensure sparsity. This reduced sample contained no angular separations in the range  $< 0.125$  deg.<sup>7</sup>

In Tables 5, 6, 7, and 8, we present the medians on the posterior distributions of our hyperparameters for our *Gaia* model, given RC-corrected seismic positions and spreads for the RC at different temperature offsets  $\Delta T_{\text{eff}}$  for both the Y18 and APOKASC-2 samples. In Fig. 7, we present the posterior distributions of  $\varpi_{\text{zp}}$  given the six values for the position of the RC used each in the *K* and *G* bands.

In order to probe the impact of literature values for  $\varpi_{\text{zp}}$  on an inference of our RC parameters, we reran our *Gaia* model for the *K* and *Gaia G* bands with a strongly informative prior on  $\varpi_{\text{zp}}$  (see equation 24). We did this for the same reduced sample of 1000 stars from our Y18 sample. For all these runs, we applied the same priors used for  $\mu_{\text{RC}}$  and  $\sigma_{\text{RC}}$  as in the asteroseismic runs (see equation 22). We used the parallax zero-point offsets reported by Lindegren et al. (2018) (−29  $\mu\text{as}$ , with an assumed uncertainty of 1  $\mu\text{as}$ ), Zinn et al. (2018) (−52.8  $\mu\text{as}$  with a total uncertainty of 3.4  $\mu\text{as}$ ), Riess et al. (2018) (−46 ± 13  $\mu\text{as}$ ), Sahlholdt & Silva Aguirre (2018) (−35 ± 16  $\mu\text{as}$ ), and Stassun & Torres (2018) (−82 ± 33  $\mu\text{as}$ ). Note that for the purpose of calibration not all these zero-point offsets would be applicable to our sample due to differences in colour, magnitude, and position. We instead used them as representative of  $\varpi_{\text{zp}}$  in the literature to study their impact on our inferences only. In addition, we also ran with a prior of  $0 \pm 1$   $\mu\text{as}$  in an attempt to re-create the H17 work (albeit accounting for parallax covariances), as well as a single run with no strongly informative priors on  $\varpi_{\text{zp}}$ ,  $\mu_{\text{RC}}$  or  $\sigma_{\text{RC}}$ , thus finding our own measure of the zero-point offset.

In Tables 9 and 10 we present the medians and  $1\sigma$  credible intervals on the posterior distributions for the hyperparameters of our *Gaia* model, given the conditions stated above, as well as naming the source of the used parallax zero-point offset, and an expression

<sup>7</sup>The data were shuffled using the `sklearn.utils.shuffle` function with a random seed of 24601.



**Figure 5.** The posterior distributions on the position of the RC in the 2MASS *K* band (left) and *Gaia G* band (right), as a function of overall perturbation to the temperature values  $\Delta T_{\text{eff}}$  using asteroseismology, both with (orange) and without (green) corrections to the  $\Delta \nu$  scaling relation (Sharma et al. 2016). The dashed horizontal lines indicate the median on the posteriors, and the solid horizontal lines represent the  $1\sigma$  credible intervals. The posteriors' magnitudes along the *x*-axis are indicative of power with arbitrary units, whereas their shape along the *y*-axis indicates the spread in the posterior result.

**Table 2.** Same as Table 1, except for a subsample of stars from the APOKASC-2 (Pinsonneault et al. 2018) sample.

$\Delta T_{\text{eff}}$ (K)	No correction				Clump corrected			
	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )
-50.0	$-1.659 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.2^{+1.2}_{-1.09}$	$-1.663 \pm 0.003$	$0.031 \pm 0.005$	$0.89 \pm 0.02$	$8.46^{+1.19}_{-1.06}$
-40.0	$-1.664 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.14^{+1.18}_{-1.08}$	$-1.669 \pm 0.003$	$0.032 \pm 0.005$	$0.89 \pm 0.02$	$8.4^{+1.16}_{-1.1}$
-30.0	$-1.669 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.13^{+1.23}_{-1.11}$	$-1.675 \pm 0.003$	$0.031 \pm 0.005$	$0.89 \pm 0.02$	$8.53^{+1.16}_{-1.06}$
-20.0	$-1.674 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.15^{+1.26}_{-1.1}$	$-1.681 \pm 0.003$	$0.031 \pm 0.005$	$0.89 \pm 0.02$	$8.43^{+1.24}_{-1.06}$
-10.0	$-1.679 \pm 0.003$	$0.03 \pm 0.004$	$0.9 \pm 0.02$	$9.11^{+1.18}_{-1.09}$	$-1.687 \pm 0.003$	$0.032 \pm 0.005$	$0.89 \pm 0.02$	$8.37^{+1.23}_{-1.11}$
0.0	$-1.684 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.13^{+1.25}_{-1.1}$	$-1.693 \pm 0.003$	$0.031 \pm 0.005$	$0.89 \pm 0.02$	$8.5^{+1.18}_{-1.08}$
10.0	$-1.689 \pm 0.003$	$0.03 \pm 0.004$	$0.9 \pm 0.02$	$9.08^{+1.2}_{-1.08}$	$-1.698 \pm 0.003$	$0.032 \pm 0.005$	$0.89 \pm 0.02$	$8.41^{+1.2}_{-1.06}$
20.0	$-1.694 \pm 0.003$	$0.03 \pm 0.004$	$0.9 \pm 0.02$	$9.04^{+1.26}_{-1.06}$	$-1.704 \pm 0.003$	$0.032 \pm 0.005$	$0.89 \pm 0.02$	$8.44^{+1.21}_{-1.08}$
30.0	$-1.699 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.1^{+1.17}_{-1.07}$	$-1.71 \pm 0.003$	$0.033 \pm 0.005$	$0.9 \pm 0.02$	$8.29^{+1.21}_{-1.05}$
40.0	$-1.704 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.12^{+1.18}_{-1.1}$	$-1.715 \pm 0.003$	$0.032 \pm 0.005$	$0.89 \pm 0.02$	$8.43^{+1.23}_{-1.07}$
50.0	$-1.709 \pm 0.003$	$0.029 \pm 0.004$	$0.9 \pm 0.02$	$9.15^{+1.24}_{-1.09}$	$-1.721 \pm 0.003$	$0.032 \pm 0.005$	$0.89 \pm 0.02$	$8.39^{+1.19}_{-1.07}$

of the prior applied to  $\varpi_{\text{zp}}$ . Note that the inferred value of  $\varpi_{\text{zp}}$  may differ significantly within the uncertainties on any of the literature values used. In Fig. 8 we present the medians and  $1\sigma$  credible intervals on the posterior distributions for  $\mu_{\text{RC}}$  given our chosen values for  $\varpi_{\text{zp}}$ , with the result from the ‘uninformed’ run shown with bold red error bars.

## 5 DISCUSSION

### 5.1 Luminosity of the Red Clump

Figs 5 and 6 show the posteriors on the inferred absolute magnitude of the RC,  $\mu_{\text{RC}}$ , for the *K* and *Gaia G* bands, given changes to effective temperature and corrections to the scaling relations. There is a clear relation between the overall offset in  $T_{\text{eff}}$  and the inferred

magnitude of the RC, where a change of about 20 K results in a difference of more than  $1\sigma$ . The overall relation between the clump magnitude and temperature is expected, given the large impact of temperature on the calculations for absolute magnitude; luminosity calculated via the seismic scaling relations scales with temperature to a power of 4.5, and BCs calculated through the Casagrande & Vandenberg (2018b) method rely on both  $T_{\text{eff}}$  and  $\log g$ , which is calculated using  $T_{\text{eff}}$ . The small uncertainties on  $\mu_{\text{RC}}$  and  $\sigma_{\text{RC}}$  indicate the ability of hierarchical models to leverage a large number of individual uncertainties to fit to a population model, given that the uncertainties on our data for  $T_{\text{eff}}$  are well above the shifts in temperature we are applying.

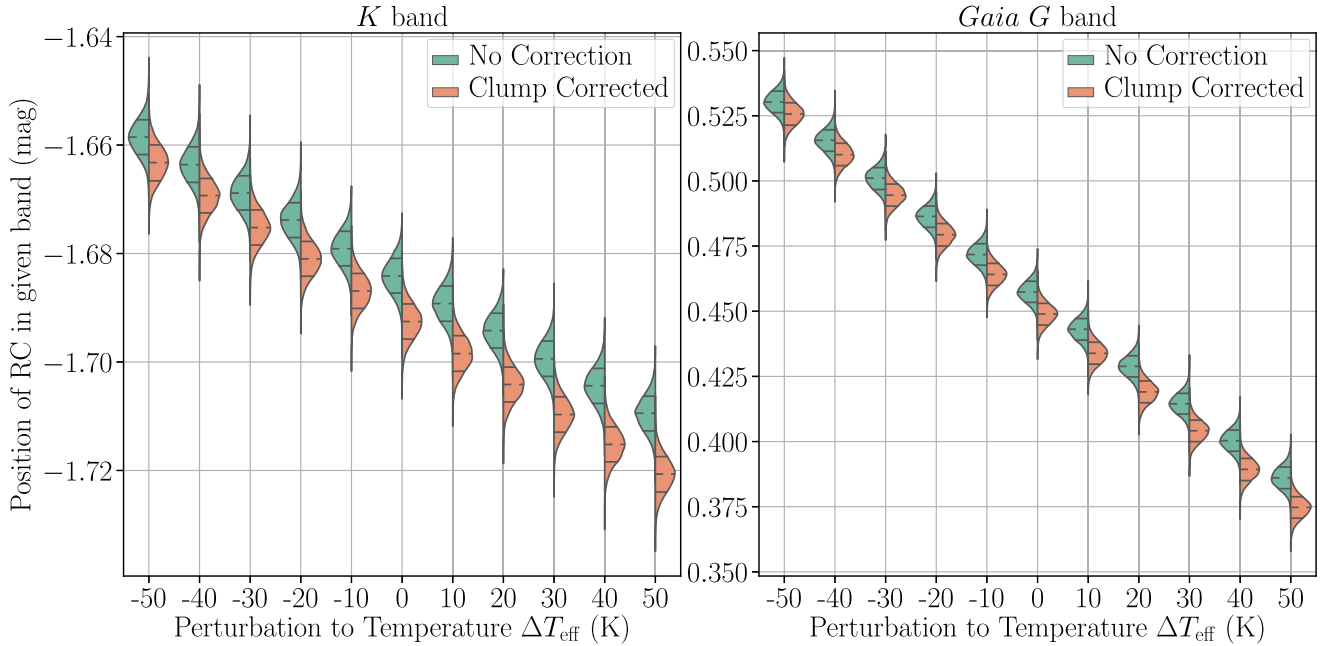
We also see that the scaling relation corrections appear to be degenerate with a small temperature offset. A change of  $\sim 20$  K to the temperatures provides a similar clump magnitude as when

**Table 3.** Same as Table 1, except for the *Gaia* *G* band, for 5576 stars from the Y18 sample.

$\Delta T_{\text{eff}}$ (K)	$\mu_{\text{RC}}$ (mag)	No correction			Clump corrected			
		$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )
−50.0	0.35 ± 0.003	0.181 ± 0.004	0.98 ± 0.01	2.73 <sup>+0.58</sup> <sub>−0.42</sub>	0.34 ± 0.003	0.193 ± 0.004	0.99 ± 0.01	2.77 <sup>+0.66</sup> <sub>−0.46</sub>
−40.0	0.336 ± 0.003	0.181 ± 0.004	0.98 ± 0.01	2.73 <sup>+0.56</sup> <sub>−0.4</sub>	0.325 ± 0.003	0.192 ± 0.004	0.99 ± 0.01	2.78 <sup>+0.66</sup> <sub>−0.45</sub>
−30.0	0.323 ± 0.003	0.18 ± 0.004	0.98 ± 0.01	2.72 <sup>+0.54</sup> <sub>−0.4</sub>	0.311 ± 0.003	0.19 ± 0.004	0.99 ± 0.01	2.79 <sup>+0.58</sup> <sub>−0.43</sub>
−20.0	0.309 ± 0.003	0.179 ± 0.004	0.98 ± 0.01	2.72 <sup>+0.54</sup> <sub>−0.4</sub>	0.297 ± 0.003	0.188 ± 0.004	0.98 ± 0.01	2.72 <sup>+0.58</sup> <sub>−0.41</sub>
−10.0	0.295 ± 0.003	0.178 ± 0.004	0.98 ± 0.01	2.69 <sup>+0.53</sup> <sub>−0.38</sub>	0.282 ± 0.003	0.187 ± 0.004	0.98 ± 0.01	2.71 <sup>+0.58</sup> <sub>−0.4</sub>
0.0	0.282 ± 0.003	0.177 ± 0.004	0.98 ± 0.01	2.68 <sup>+0.52</sup> <sub>−0.38</sub>	0.268 ± 0.003	0.187 ± 0.004	0.98 ± 0.01	2.73 <sup>+0.58</sup> <sub>−0.42</sub>
10.0	0.268 ± 0.003	0.177 ± 0.004	0.98 ± 0.01	2.71 <sup>+0.53</sup> <sub>−0.39</sub>	0.254 ± 0.003	0.185 ± 0.004	0.98 ± 0.01	2.7 <sup>+0.58</sup> <sub>−0.41</sub>
20.0	0.255 ± 0.003	0.176 ± 0.004	0.98 ± 0.01	2.7 <sup>+0.51</sup> <sub>−0.37</sub>	0.24 ± 0.003	0.184 ± 0.004	0.98 ± 0.01	2.71 <sup>+0.56</sup> <sub>−0.4</sub>
30.0	0.241 ± 0.003	0.175 ± 0.004	0.98 ± 0.01	2.68 <sup>+0.51</sup> <sub>−0.36</sub>	0.226 ± 0.003	0.183 ± 0.004	0.98 ± 0.01	2.7 <sup>+0.55</sup> <sub>−0.4</sub>
40.0	0.228 ± 0.003	0.174 ± 0.004	0.98 ± 0.01	2.67 <sup>+0.48</sup> <sub>−0.36</sub>	0.213 ± 0.003	0.182 ± 0.004	0.98 ± 0.01	2.7 <sup>+0.52</sup> <sub>−0.4</sub>
50.0	0.215 ± 0.003	0.173 ± 0.004	0.98 ± 0.01	2.68 <sup>+0.48</sup> <sub>−0.36</sub>	0.199 ± 0.003	0.181 ± 0.004	0.98 ± 0.01	2.69 <sup>+0.53</sup> <sub>−0.38</sub>

**Table 4.** Same as Table 1, except for the *Gaia* *G* Band, for a subsample of stars from the APOKASC-2 (Pinsonneault et al. 2018) sample.

$\Delta T_{\text{eff}}$ (K)	$\mu_{\text{RC}}$ (mag)	No correction			Clump corrected			
		$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )
−50.0	0.53 ± 0.004	0.118 ± 0.006	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.19 <sup>+0.65</sup> <sub>−0.46</sub>	0.526 ± 0.004	0.128 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.29 <sup>+0.81</sup> <sub>−0.55</sub>
−40.0	0.516 ± 0.004	0.117 ± 0.005	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.18 <sup>+0.65</sup> <sub>−0.46</sub>	0.51 ± 0.004	0.127 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.31 <sup>+0.8</sup> <sub>−0.55</sub>
−30.0	0.501 ± 0.004	0.116 ± 0.006	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.19 <sup>+0.62</sup> <sub>−0.45</sub>	0.495 ± 0.004	0.127 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.29 <sup>+0.76</sup> <sub>−0.54</sub>
−20.0	0.486 ± 0.004	0.116 ± 0.006	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.19 <sup>+0.65</sup> <sub>−0.47</sub>	0.479 ± 0.004	0.126 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.28 <sup>+0.77</sup> <sub>−0.54</sub>
−10.0	0.472 ± 0.004	0.115 ± 0.006	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.19 <sup>+0.63</sup> <sub>−0.45</sub>	0.464 ± 0.004	0.126 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.29 <sup>+0.74</sup> <sub>−0.55</sub>
0.0	0.457 ± 0.004	0.114 ± 0.006	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.2 <sup>+0.64</sup> <sub>−0.45</sub>	0.449 ± 0.004	0.125 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.27 <sup>+0.79</sup> <sub>−0.53</sub>
10.0	0.443 ± 0.004	0.113 ± 0.006	0.95 <sup>+0.02</sup> <sub>−0.03</sub>	3.17 <sup>+0.63</sup> <sub>−0.44</sub>	0.434 ± 0.004	0.124 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.25 <sup>+0.8</sup> <sub>−0.52</sub>
20.0	0.429 ± 0.004	0.113 ± 0.006	0.96 <sup>+0.02</sup> <sub>−0.03</sub>	3.21 <sup>+0.62</sup> <sub>−0.44</sub>	0.419 ± 0.004	0.124 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.25 <sup>+0.73</sup> <sub>−0.53</sub>
30.0	0.414 ± 0.004	0.112 ± 0.006	0.95 <sup>+0.02</sup> <sub>−0.03</sub>	3.18 <sup>+0.61</sup> <sub>−0.43</sub>	0.404 ± 0.004	0.123 ± 0.005	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.25 <sup>+0.72</sup> <sub>−0.54</sub>
40.0	0.4 ± 0.004	0.112 ± 0.006	0.95 <sup>+0.02</sup> <sub>−0.03</sub>	3.19 <sup>+0.59</sup> <sub>−0.44</sub>	0.389 ± 0.004	0.122 ± 0.005	0.97 <sup>+0.02</sup> <sub>−0.02</sub>	3.24 <sup>+0.72</sup> <sub>−0.5</sub>
50.0	0.386 ± 0.004	0.111 ± 0.006	0.95 <sup>+0.02</sup> <sub>−0.03</sub>	3.2 <sup>+0.57</sup> <sub>−0.43</sub>	0.375 ± 0.004	0.122 ± 0.006	0.97 <sup>+0.01</sup> <sub>−0.02</sub>	3.25 <sup>+0.71</sup> <sub>−0.51</sub>

**Figure 6.** The same as Fig. 5, except using only stars both in our sample and in the APOKASC-2 (Pinsonneault et al. 2018) sample, replacing  $T_{\text{eff}}$  with those reported in APOKASC-2.

**Table 5.** Medians of the posterior distributions for hyperparameters of our *Gaia* model, for the 2MASS *K* band, for a randomly selected subsample of 1000 stars from the Y18 sample. Uncertainties are taken as the  $1\sigma$  credible intervals, and are listed as a single value for cases where the posterior was approximately Gaussian. Priors were imposed on  $\mu_{RC}$  and  $\sigma_{RC}$  corresponding to the results for these values using seismic Clump Corrected data in Table 1, for the temperature shifts shown in the  $\Delta T_{\text{eff}}$  column.  $L$  is the length scale of the exponentially decaying space density prior on distance (Bailer-Jones et al. 2018), and  $\varpi_{zp}$  is the parallax zero-point offset. All other symbols are the same as for Table 1.

$\Delta T_{\text{eff}}$ (K)	$\mu_{RC}$ (mag)	$\sigma_{RC}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{RC}$ )	$L$ (pc)	$\varpi_{zp}$ ( $\mu\text{as}$ )
−50.0	$-1.71 \pm 0.002$	$0.041 \pm 0.003$	$0.58^{+0.05}_{-0.05}$	$5.49^{+0.52}_{-0.47}$	$908.63^{+16.55}_{-15.89}$	$-24.09^{+12.84}_{-12.76}$
0.0	$-1.737 \pm 0.002$	$0.04 \pm 0.003$	$0.55^{+0.05}_{-0.03}$	$5.61^{+0.5}_{-0.47}$	$920.12^{+17.18}_{-16.61}$	$-19.5^{+12.4}_{-12.46}$
50.0	$-1.764 \pm 0.002$	$0.041 \pm 0.004$	$0.53^{+0.04}_{-0.02}$	$5.5^{+0.5}_{-0.48}$	$930.95^{+18.07}_{-16.83}$	$-14.81^{+12.57}_{-12.98}$

**Table 6.** Same as Table 5, except for priors imposed on  $\mu_{RC}$  and  $\sigma_{RC}$  corresponding to the results for these values using seismic Clump Corrected data in Table 2, for the temperature shifts shown in the  $\Delta T_{\text{eff}}$  column.

$\Delta T_{\text{eff}}$ (K)	$\mu_{RC}$ (mag)	$\sigma_{RC}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{RC}$ )	$L$ (pc)	$\varpi_{zp}$ ( $\mu\text{as}$ )
−50.0	$-1.661 \pm 0.003$	$0.04 \pm 0.003$	$0.6 \pm 0.05$	$5.76^{+0.55}_{-0.5}$	$888^{+16.38}_{-15.78}$	$-33.53^{+12.93}_{-12.97}$
0.0	$-1.689 \pm 0.003$	$0.04 \pm 0.004$	$0.59 \pm 0.05$	$5.66^{+0.53}_{-0.51}$	$899.36^{+16.72}_{-16.04}$	$-28.33^{+12.96}_{-12.92}$
50.0	$-1.715 \pm 0.003$	$0.041 \pm 0.004$	$0.57 \pm 0.05$	$5.51^{+0.55}_{-0.49}$	$910.68^{+16.83}_{-16.4}$	$-23.47^{+13.25}_{-13.13}$

**Table 7.** Same as Table 5, except for the *Gaia* *G* band, with priors imposed on  $\mu_{RC}$  and  $\sigma_{RC}$  corresponding to the results for these values using seismic Clump Corrected data in Table 3, for the temperature shifts shown in the  $\Delta T_{\text{eff}}$  column.

$\Delta T_{\text{eff}}$ (K)	$\mu_{RC}$ (mag)	$\sigma_{RC}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{RC}$ )	$L$ (pc)	$\varpi_{zp}$ ( $\mu\text{as}$ )
−50.0	$0.346 \pm 0.003$	$0.19 \pm 0.003$	$0.97^{+0.01}_{-0.02}$	$3.1^{+0.79}_{-0.7}$	$948.41^{+18.15}_{-17.96}$	$-9.96^{+13.1}_{-13.18}$
0.0	$0.277 \pm 0.003$	$0.188 \pm 0.004$	$0.95^{+0.02}_{-0.05}$	$2.64^{+0.77}_{-0.63}$	$978.9^{+18.35}_{-17.46}$	$1.14^{+12.8}_{-12.81}$
50.0	$0.209 \pm 0.003$	$0.184 \pm 0.004$	$0.74^{+0.12}_{-0.13}$	$1.71^{+0.36}_{-0.2}$	$1008.77^{+18.85}_{-18.48}$	$10.76^{+13.13}_{-13.21}$

**Table 8.** Same as Table 5, except for the *Gaia* *G* band, with priors imposed on  $\mu_{RC}$  and  $\sigma_{RC}$  corresponding to the results for these values using seismic Clump Corrected data in Table 4, for the temperature shifts shown in the  $\Delta T_{\text{eff}}$  column.

$\Delta T_{\text{eff}}$ (K)	$\mu_{RC}$ (mag)	$\sigma_{RC}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{RC}$ )	$L$ (pc)	$\varpi_{zp}$ ( $\mu\text{as}$ )
−50.0	$0.527 \pm 0.004$	$0.13 \pm 0.005$	$0.82^{+0.05}_{-0.07}$	$2.53^{+0.36}_{-0.3}$	$874.12^{+16.56}_{-16.1}$	$-39.02^{+12.98}_{-13.16}$
0.0	$0.455 \pm 0.004$	$0.129 \pm 0.005$	$0.79^{+0.07}_{-0.08}$	$2.42^{+0.36}_{-0.29}$	$903.23^{+16.8}_{-16.68}$	$-26.84^{+13.1}_{-12.97}$
50.0	$0.385 \pm 0.004$	$0.127 \pm 0.005$	$0.68^{+0.09}_{-0.1}$	$2.22^{+0.27}_{-0.22}$	$931.92^{+17.53}_{-17}$	$-14.94^{+12.58}_{-13.04}$

applying a correction to the scaling relations. At higher temperatures, the difference in the magnitude of the RC between corrected and uncorrected scaling relations increases. This shows that the  $T_{\text{eff}}$  values have a significant impact on the  $f_{\Delta v}$  obtained through the Sharma & Stello (2016) method, even at relatively small  $T_{\text{eff}}$  shifts.

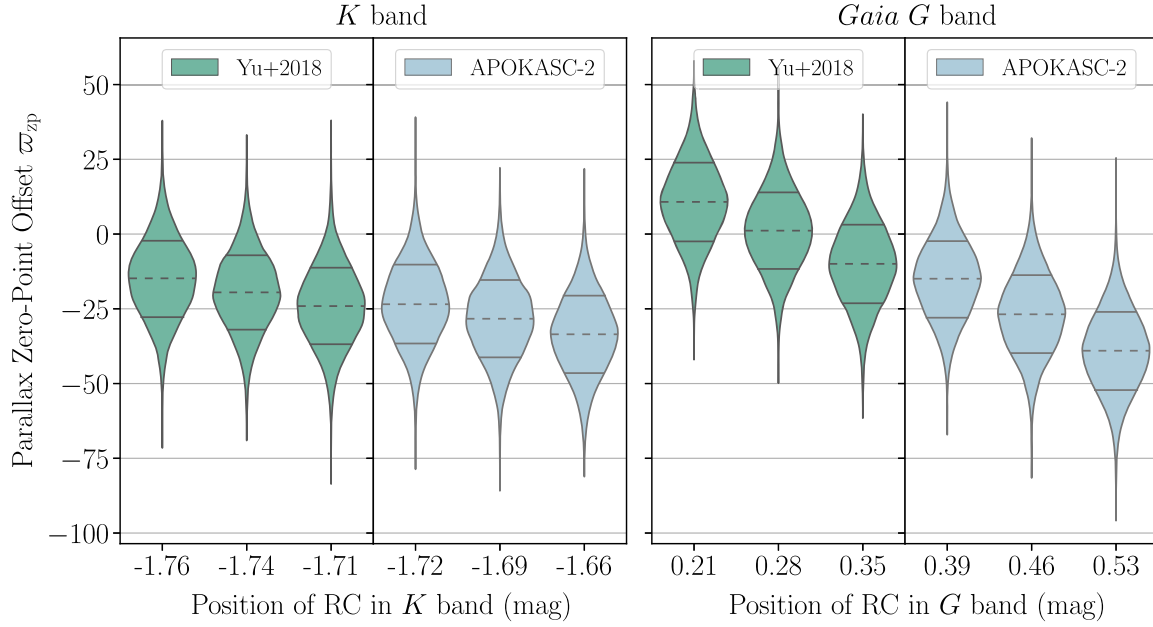
The values for  $\mu_{RC}$  in both bands are fainter for the subset of stars using APOKASC-2 temperatures than those using temperatures from Mathur et al. (2017). This reflects the relation we already saw between  $T_{\text{eff}}$  and  $\mu_{RC}$  for the Y18 stars, since the stars in the APOKASC-2 subsample represent a population subset of lower-temperature stars, as well as having lower values for  $T_{\text{eff}}$  in the APOKASC-2 catalogue itself. However, the fact that APOKASC-2 stars represent a lower-temperature population only accounts for a shift in a measured median absolute magnitude of  $\sim 0.028$  mag in *K* and  $\sim 0.12$  mag in *G*. The use of APOKASC-2 temperatures for the subset shifts the absolute magnitudes even fainter, by another  $\sim 0.028$  mag and  $\sim 0.07$  mag in *K* and *G*, respectively. At the precision afforded to us by hierarchical models, these shifts caused by the choice of temperatures become statistically significant.

Due to the nature of the *K* band minimizing the effects of metallicity on the RC spread, there is an extensive literature on the

value of  $\mu_{RC}$  in *K*. It was found by Alves (2000) to be  $-1.62 \pm 0.03$  (with a consistent measurement by Udalski 2000), but later placed at  $-1.54 \pm 0.04$  by Groenewegen (2008). A recent review by Girardi (2016) found a median literature value of  $-1.59 \pm 0.04$  mag, which was applied by Davies et al. (2017) to calibrate TGAS parallaxes. New work by Chen et al. (2017) has used RC stars similarly identified using asteroseismology to find  $-1.626 \pm 0.057$  mag, and the precursor to our hierarchical Bayesian approach, H17, used TGAS parallaxes to find  $-1.61 \pm 0.01$  mag. Using the same method, H17 reported an absolute magnitude of  $0.44 \pm 0.01$  mag in the *Gaia* *G* band.

Our RC magnitudes for both the *K* and *Gaia* *G* bands are much closer to those reported in literature when we used APOKASC-2 stars and temperatures alone. For the *K* band, we found values within  $1\sigma$  of Chen et al. (2017) for  $\Delta T_{\text{eff}} \leq 20$  K when using corrections to the scaling relations, although our results are otherwise incompatible with the literature for *K*. In the *G* band, however, we found values for  $\mu_{RC}$  compatible with H17 when using APOKASC-2 stars for  $\Delta T_{\text{eff}}$  of 0 or  $+10$  K both with and without corrections to the scaling relations. The disagreement found only in the *K* band could be due to our choice of BCs or corrections to the scaling





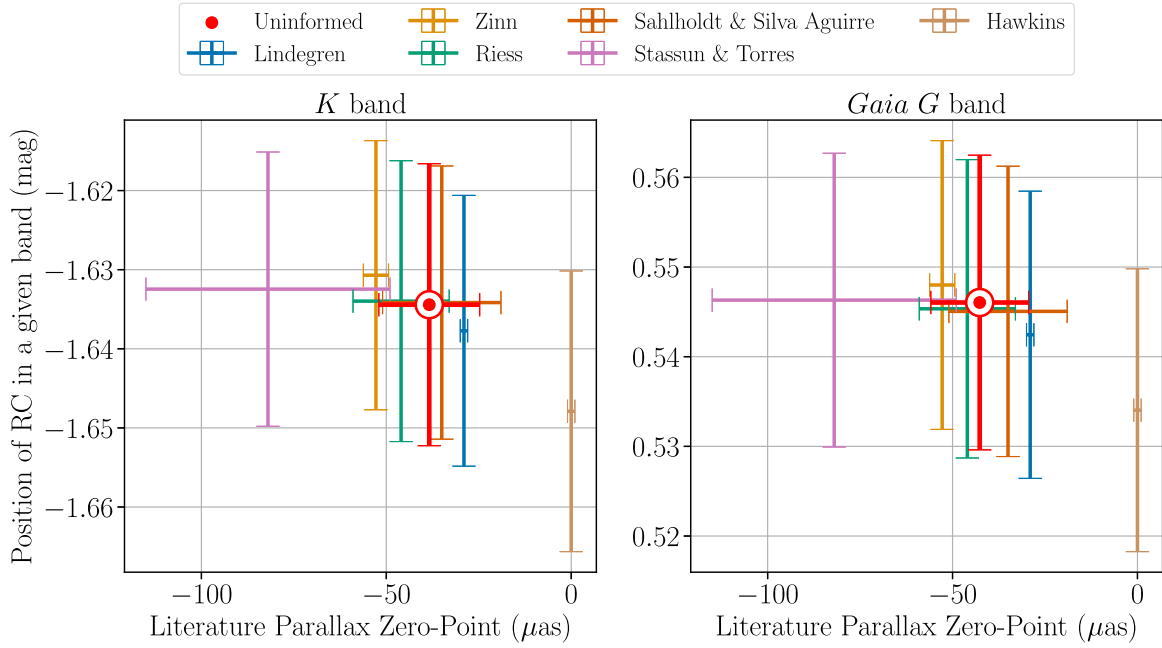
**Figure 7.** The posterior distributions on the parallax zero-point offset  $\varpi_{zp}$ , as a function of the absolute magnitude of the RC used to calibrate this value, for 1000 randomly selected stars across the *Kepler* field. The RC magnitudes on the x-axis correspond to those obtained from seismology for perturbations to the temperature values  $\Delta T_{\text{eff}}$  of  $-50$ ,  $0$ , and  $+50$  K, from runs on our full sample (Yu et al. 2018) and the APOKASC-2 sample (Pinsonneault et al. 2018). The dashed horizontal lines indicate the median on the posteriors, and the solid horizontal lines represent the  $1\sigma$  credible intervals. The posteriors' magnitudes along the x-axis are indicative of power with arbitrary units, whereas their shape along the y-axis indicates the spread in the posterior result, and is reflected across the x-axis.

**Table 9.** Medians on the posterior distributions for hyperparameters on our *Gaia* model, for the 2MASS *K* band, for a randomly selected subsample of 1000 stars from the Y18 sample. Uncertainties are taken as the  $1\sigma$  credible intervals, and are listed as single values for cases where the posterior was approximately Gaussian. Highly informative priors, shown in the ' $\varpi_{zp}$  prior' column, were imposed on  $\varpi_{zp}$  corresponding to estimates for this parameter from the literature, listed in bold print in the Source column. Additionally, we apply a custom prior to place  $\varpi_{zp}$  near zero in order to re-create conditions similar to the H17 work, and an extremely broad prior on  $\varpi_{zp}$  in order to find a value given no strong constraints on neither  $\varpi_{zp}$ ,  $\mu_{\text{RC}}$  nor  $\sigma_{\text{RC}}$ .  $\mathcal{N}(\mu, \sigma)$  indicates a normal distribution with mean  $\mu$  and standard deviation  $\sigma$ .

Source	$\varpi_{zp}$ prior ( $\mu\text{as}$ )	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )	$L$ (pc)	$\varpi_{zp}$ ( $\mu\text{as}$ )
<b>Lindgren + 18</b>	$\mathcal{N}(-29.0, 1.0)$	$-1.638 \pm 0.017$	$0.075^{+0.016}_{-0.015}$	$0.78^{+0.09}_{-0.11}$	$3.28^{+0.64}_{-0.56}$	$888.56^{+25.41}_{-24.36}$	$-29.07^{+1}_{-0.99}$
<b>Zinn + 18</b>	$\mathcal{N}(-52.8, 3.4)$	$-1.631 \pm 0.017$	$0.074^{+0.016}_{-0.015}$	$0.77^{+0.09}_{-0.1}$	$3.3^{+0.65}_{-0.57}$	$885.76^{+24.4}_{-23.73}$	$-51.92^{+3.21}_{-3.21}$
<b>Riess + 18</b>	$\mathcal{N}(-46.0, 13.0)$	$-1.634 \pm 0.017$	$0.076^{+0.017}_{-0.015}$	$0.78^{+0.09}_{-0.11}$	$3.26^{+0.64}_{-0.57}$	$886.59^{+25}_{-24.12}$	$-42.22^{+9.33}_{-9.16}$
<b>Sahlholdt &amp; Silva Aguirre 18</b>	$\mathcal{N}(-35.0, 16.0)$	$-1.634 \pm 0.017$	$0.073^{+0.016}_{-0.015}$	$0.77^{+0.09}_{-0.11}$	$3.33^{+0.69}_{-0.58}$	$887.37^{+24.06}_{-23.89}$	$-37^{+10.17}_{-10.42}$
<b>Stassun &amp; Torres 18</b>	$\mathcal{N}(-82.0, 33.0)$	$-1.632 \pm 0.017$	$0.072^{+0.017}_{-0.016}$	$0.76^{+0.09}_{-0.11}$	$3.36^{+0.64}_{-0.59}$	$885.77^{+24.33}_{-23.01}$	$-44.55^{+12.62}_{-12.59}$
<i>Hawkins + 17</i>	$\mathcal{N}(0.0, 1.0)$	$-1.648 \pm 0.018$	$0.075^{+0.017}_{-0.015}$	$0.78^{+0.09}_{-0.11}$	$3.31^{+0.64}_{-0.57}$	$893.39^{+24.6}_{-24}$	$-0.22^{+0.99}_{-1.01}$
<i>Uninformed</i>	$\mathcal{N}(0.0, 1000.0)$	$-1.634 \pm 0.018$	$0.074^{+0.017}_{-0.015}$	$0.77^{+0.09}_{-0.11}$	$3.3^{+0.64}_{-0.58}$	$887.27^{+24.12}_{-23.82}$	$-38.38^{+13.83}_{-13.54}$

**Table 10.** Same as Table 9, except for the *Gaia* *G* band.

Source	$\varpi_{zp}$ prior ( $\mu\text{as}$ )	$\mu_{\text{RC}}$ (mag)	$\sigma_{\text{RC}}$ (mag)	$Q$	$\sigma_o$ ( $\sigma_{\text{RC}}$ )	$L$ (pc)	$\varpi_{zp}$ ( $\mu\text{as}$ )
<b>Lindgren + 18</b>	$\mathcal{N}(-29.0, 1.0)$	$0.542 \pm 0.016$	$0.138^{+0.014}_{-0.018}$	$0.86^{+0.07}_{-0.12}$	$2.61^{+0.48}_{-0.34}$	$868.2^{+17.41}_{-17.09}$	$-29.06^{+0.98}_{-1.01}$
<b>Zinn + 18</b>	$\mathcal{N}(-52.8, 3.4)$	$0.548 \pm 0.016$	$0.139^{+0.014}_{-0.018}$	$0.86^{+0.07}_{-0.12}$	$2.62^{+0.49}_{-0.35}$	$865.44^{+16.95}_{-17.15}$	$-52.18^{+3.27}_{-3.31}$
<b>Riess + 18</b>	$\mathcal{N}(-46.0, 13.0)$	$0.545 \pm 0.016$	$0.14^{+0.013}_{-0.017}$	$0.87^{+0.07}_{-0.11}$	$2.62^{+0.48}_{-0.34}$	$867.13^{+17.23}_{-17.55}$	$-44.23^{+9.06}_{-9.32}$
<b>Sahlholdt &amp; Silva Aguirre 18</b>	$\mathcal{N}(-35.0, 16.0)$	$0.545 \pm 0.016$	$0.136^{+0.015}_{-0.021}$	$0.85^{+0.08}_{-0.14}$	$2.62^{+0.47}_{-0.34}$	$867.15^{+17.3}_{-17.05}$	$-39.29^{+9.86}_{-10.27}$
<b>Stassun &amp; Torres 18</b>	$\mathcal{N}(-82.0, 33.0)$	$0.546 \pm 0.017$	$0.138^{+0.014}_{-0.018}$	$0.86^{+0.07}_{-0.12}$	$2.61^{+0.46}_{-0.33}$	$866.11^{+17.76}_{-17.02}$	$-47.86^{+12.18}_{-12.51}$
<i>Hawkins + 17</i>	$\mathcal{N}(0.0, 1.0)$	$0.534 \pm 0.015$	$0.14^{+0.013}_{-0.018}$	$0.87^{+0.06}_{-0.12}$	$2.64^{+0.5}_{-0.35}$	$872.01^{+17.8}_{-17.38}$	$-0.23^{+1}_{-1.01}$
<i>Uninformed</i>	$\mathcal{N}(0.0, 1000.0)$	$0.546 \pm 0.016$	$0.139^{+0.013}_{-0.019}$	$0.87^{+0.07}_{-0.13}$	$2.62^{+0.49}_{-0.34}$	$866.26^{+17.53}_{-16.86}$	$-42.66^{+13.48}_{-13.14}$



**Figure 8.** The  $1\sigma$  credible intervals for the posterior distributions on  $\mu_{RC}$ , as a function of the value for  $\varpi_{zp}$  used as an informative prior on  $\varpi_{zp}$ , for 1000 randomly selected stars across the *Kepler* field in both the 2MASS *K* and *Gaia* *G* bands. The errorbars on the *x*-axis correspond to the formal uncertainties for literature values, or are otherwise specified in the text. The ‘uninformed’ value corresponds to a run of our *Gaia* model with no strong constraints on  $\varpi_{zp}$ , and in this case the *x*-axis errorbars correspond to the  $1\sigma$  credible intervals on the inferred value for  $\varpi_{zp}$ .

relations, or due to H17’s choice of extinction coefficient, which is twice as large as the coefficient we use in our *Gaia* models, and would bias the absolute magnitudes of their stars towards brighter values. Alternatively, it could be due to H17 not accounting for known spatial correlations in parallax (Lindegren et al. 2016; Zinn et al. 2017) or possible parallax zero-point offsets (Brown 2018).

In Tables 9 and 10, we attempt to re-create the H17 work, albeit including parallax covariances, and find values for  $\mu_{RC}$  that are compatible with a temperature offset of  $\Delta T_{\text{eff}} < -50$  K for both photometric bands. Finally, allowing  $\varpi_{zp}$  to vary as a free parameter with loose prior constraints finds  $\mu_{RC} = -1.634 \pm 0.018$  mag in the *K* band and  $0.546 \pm 0.016$  mag in the *G* band. These values imply that a shift to the temperature scales of  $-50$  K or more is appropriate when using temperatures for seismology of the RC.

## 5.2 Spread of the Red Clump

In principle, the spread of the RC, like its luminosity, is a property of an RC population and depends on the mass and metallicity of the sample (Salaris & Girardi 2002; Girardi 2016). Our hierarchical approach allows us to study the ‘true’ spread of the RC, by evaluating the uncertainties on individual measures of absolute magnitude.

As seen for the *K* band in Tables 1 and 2, the spread of the RC is consistent within  $1\sigma$  for all perturbations of temperature, corrections to the scaling relations, and between both the Yu et al. (2018) and APOKASC-2 temperatures. This indicates that  $\sigma_{RC}$  is only weakly dependent on the choice of temperature scale, and that any effects of the APOKASC-2 sample only representing a small subset in metallicity are minimal for the *K* band.

The spread of the RC due to mass and metallicity is minimized in the 2MASS *K* band (Salaris & Girardi 2002), which would lead us to expect a broader spread of the RC in the *Gaia* *G* band. We see this effect in Tables 3 and 4, where the reported spreads are

$\sim 4$  to 6 times larger in magnitude. Surprisingly, we do not see the same consistency for the values of  $\sigma_{RC}$  for the *G* band, but instead find that the inferred value of  $\sigma_{RC}$  varies inversely with temperature beyond  $1\sigma$  from  $-50$  K to  $50$  K. This trend of  $\sigma_{RC}$  with  $\Delta T_{\text{eff}}$  is likely to be an effect of the BC, as we do not see a compatible trend in *K*. It should also be noted that we would expect extinction to play a larger role in the *G* band, possibly contributing to this effect.

For the *Gaia* *G* band we also see that the value for  $\sigma_{RC}$  is lower for the APOKASC-2 sample than for the full Y18 sample. This reduction is likely because the APOKASC-2 sample draws temperatures from a uniform spectroscopy source (and thus temperature scale) whereas the Y18 temperatures come from a variety of sources, broadening the distribution of RC stars.

The similar hierarchical approach taken by H17 found a spread of  $0.17 \pm 0.02$  mag in *K* and  $0.20 \pm 0.02$  mag in *G* using TGAS parallaxes. The agreement within  $1\sigma$  for the *G* band for the Y18 sample agrees with the inferred APOKASC-2 spread being an underestimate. The estimates found in our work for  $\sigma_{RC}$  in *K* are an order of magnitude lower. This is probably due to our sample size (increased from H17 by a factor of 5) and asteroseismology providing more precise measurements for these stars than TGAS (Davies et al. 2017), allowing the hierarchical method to more closely constrain the true underlying spread.

Tables 9 and 10 show the results of our attempt to re-create the H17 work, accounting for parallax covariances and including a parallax zero-point offset. Using *Gaia* parallaxes, we found an  $\sigma_{RC}$  in *K* that is larger than our value from seismology. The results presented in Tables 5 and 6, where the seismic  $\sigma_{RC}$  in *K* has been applied as a prior on the *Gaia* model, show an inlier fraction  $Q$  that is lower than we would expect for this sample. This implies that *Gaia* DR2 is underestimating the uncertainties for stars considered ‘outliers’, and not including them in the inlier population.

For the *G* band, we found a value for  $\sigma_{RC}$  in agreement with our seismic value using APOKASC-2 temperatures. In this instance, as

opposed to the results shown in Table 8 at similar  $\sigma_{\text{RC}}$ , we find an inlier fraction  $Q$  in the expected range. This is probably due to the simultaneous inference of a more appropriate value for  $\mu_{\text{RC}}$ , which is closer to values established in literature (H17). For this reason, the spreads reported in Tables 4 and 10 are our best estimates for the ‘true’ spread of the RC in the  $G$  band.

With our measurement of  $\sigma_{\text{RC}} = 0.03$  mag in the  $K$  band, we can use standard error propagation through equation 20 (setting extinction to zero) to find that this spread yields a precision in distance of  $\sim 1$  per cent for our sample, subject to mass and metallicity. This is a factor of 5 improvement from the precision reported by H17. When using  $\sigma_{\text{RC}} = 0.14$  mag for the  $G$  band, we find a distance precision of  $\sim 6$  per cent, in line with the findings by H17.

### 5.3 The *Gaia* parallax zero-point offset

The *Gaia* DR2 parallax zero-point offset, while small, can still have an effect on results, and is widely applied in studies using DR2 (Bailer-Jones et al. 2018; Luri et al. 2018), with potentially far-reaching consequences (Shanks, Hogarth & Metcalfe 2019). The offset has been estimated through calibration with eclipsing binaries (Stassun & Torres 2018), Cepheids (Riess et al. 2018), asteroseismology (Sahlholdt & Silva Aguirre 2018; Zinn et al. 2018), kinematics (Schönrich, McMillan & Eyer 2019), and quasars (Lindgren et al. 2018).

In Tables 5, 6, 7, and 8, we present our inferred model parameters given our values for  $\mu_{\text{RC}}$  and  $\sigma_{\text{RC}}$  found through asteroseismology at different temperature shifts  $\Delta T_{\text{eff}}$ , effectively ‘calibrating’ *Gaia* DR2 to see what offset recovers a given set of RC parameters.

Fig. 7 shows the posterior distributions for  $\varpi_{\text{zp}}$  given our seismic priors from different temperature shifts, where there is a clear trend of  $\varpi_{\text{zp}}$  with seismic  $\mu_{\text{RC}}$ , and thus with temperature. This trend was also found in recent results by Khan et al. (2019), where a comparison of *Gaia* parallaxes and seismic distances obtained through the seismic scaling relations found that a temperature shift of 100 K caused a shift in  $\varpi_{\text{zp}}$  of 10 – 15  $\mu\text{as}$  for RC stars, although it should be noted that they found this effect largely reduced when using grid modelling techniques (Rodrigues et al. 2017).

It is also apparent in Fig. 7 that the uncertainty on  $\varpi_{\text{zp}}$  is significant, and consistent for all model conditions, due to the parallax covariances presenting a systematic lower limit on parallax uncertainties for this sample. Given a  $\mu_{\text{RC}}$  in  $K$  closer to literature values, with the run corresponding to APOKASC-2 temperatures using  $\Delta T_{\text{eff}} = -50$  K, we found a  $\varpi_{\text{zp}}$  within  $1\sigma$  of the uncertainties on all literature values for  $\varpi_{\text{zp}}$  in the *Kepler* field discussed in this work. This is both an encouraging sign of a consistent  $\varpi_{\text{zp}}$  in the *Kepler* field, and further indication that seismology would be improved by reducing the temperature scale. For the *Gaia*  $G$  band, the run closest to the existing literature ( $\Delta T_{\text{eff}} = 0$ ) is consistent with all values for  $\varpi_{\text{zp}}$  besides Stassun & Torres (2018).

Given a selection of values for  $\varpi_{\text{zp}}$  reported in the literature, we applied informative priors on  $\varpi_{\text{zp}}$  in our *Gaia* model, and allowed  $\mu_{\text{RC}}$  and  $\sigma_{\text{RC}}$  to explore the parameter space freely. The results of this are shown in Tables 9 and 10, for the  $K$  and  $G$  bands, respectively. The credible intervals for  $\mu_{\text{RC}}$  are shown in Fig. 8. For both bands, we found that the choice of  $\varpi_{\text{zp}}$  from the literature had no impact beyond  $1\sigma$  on either of the RC properties for any values used. When using a tightly constrained  $\varpi_{\text{zp}}$  of zero (in an attempt to re-create H17) we found the largest overall change. It is also interesting to note that for a prior corresponding to the Stassun & Torres (2018) value, the inferred value for  $\varpi_{\text{zp}}$  is reduced to lie closer to those found in other works for the *Kepler* field.

Finally, running the *Gaia* model with uninformative priors on both  $\varpi_{\text{zp}}$  and the RC parameters produced a parallax zero-point offset of  $(-38 \pm 13) \mu\text{as}$  in  $K$  and  $(-42 \pm 13) \mu\text{as}$  in  $G$  for the *Kepler* field. These values are consistent with one another and with the existing literature, and also agree with recent results by Khan et al. (2019) for RC stars in APOKASC-2. Given the uncertainties on the inferred values of  $\varpi_{\text{zp}}$ , we see a fundamental uncertainty limit on *Gaia* parallaxes of  $\sim 13 \mu\text{as}$  as a result of spatial covariances in the parallaxes. Encouragingly, this implies that for our RC sample in the *Kepler* field, the choice of parallax-zero-point offset does not dramatically impact the inferred luminosities, given a proper treatment of the spatial parallax covariances. However, this may not generalize to populations more sparsely sampled in space, and in other magnitude ranges, given the known relation between the parallax zero-point offset,  $G$ -band magnitude, and colour (Lindgren et al. 2018; Zinn et al. 2018).

### 5.4 Corrections to the seismic scaling relations

In Section 5.3, we have compared results with and without corrections to the  $\Delta\nu$  seismic scaling relation,  $f_{\Delta\nu}$ , derived from Sharma & Stello (2016). It is known that stellar models do not accurately reproduce the  $\Delta\nu$  of the Sun (off by about 1 per cent), due to the so-called surface effect (Christensen-Dalsgaard, Dappen & Lebreton 1988; White et al. 2011). Corrections to the scaling relation  $f_{\Delta\nu}$  derived without accounting for the surface effect (such as Sharma & Stello 2016) can produce radii that differ on the order of  $\sim 2$  per cent from methods that do (such as Rodrigues et al. 2017). As a check, we considered the impact that this may have on our inferred values for the RC magnitude.

To compare the calculated RC populations in the  $K$  and *Gaia*  $G$  bands, we obtained radii, using  $f_{\Delta\nu}$  obtained through Sharma & Stello (2016). We then used BCs for no temperature offset to calculate the absolute magnitudes using both those radii and those same radii reduced by both 1.6 per cent and 2.4 per cent. We found that a reduction on radius in the range  $(2 \pm 0.4)$  per cent resulted in a global shift toward brighter bolometric magnitudes by  $44^{+9}_{-8}$  mmag.

In Tables 2 and 4, we report the absolute magnitude of the RC (for no temperature offset) in the APOKASC-2,  $f_{\Delta\nu}$ -corrected, sample of  $-1.69$  mag in  $K$  and  $0.45$  mag in  $G$ . A shift of  $0.04$  mag applied to both bands is enough to reconcile our seismic results with those obtained through *Gaia* for both the  $K$  and  $G$  bands, as well as those from the literature. Note however that this is not the case when applied to the Y18 sample (see Tables 1 and 3), where this shift applied in both bands would not be enough to reconcile the seismic results for the absolute magnitude of the RC with any measures both in this work and in the literature.

### 5.5 Calibrating *Gaia* and asteroseismology

Our initial aim with this work was to calibrate the *Gaia* parallax zero-point offset,  $\varpi_{\text{zp}}$ , using asteroseismology. Given the large change in the absolute magnitude of the RC,  $\mu_{\text{RC}}$ , with relatively small changes in temperature for our large RC population, and consequently the shift in inferred  $\varpi_{\text{zp}}$  given these values for  $\mu_{\text{RC}}$ , it proved difficult to definitively calibrate *Gaia* parallaxes using seismology.

The reverse, however, seems more possible. We found that the various parallax offsets reported in the literature, when used as informative priors on our *Gaia* model, all resulted in similar values for  $\mu_{\text{RC}}$  in both the 2MASS  $K$  and *Gaia*  $G$  bands (as shown in Tables 9 and 10), and inferred values for  $\varpi_{\text{zp}}$  that lie closer together

for those literature values with large uncertainties (Riess et al. 2018; Sahlholdt & Silva Aguirre 2018; Stassun & Torres 2018). Imposing a prior for  $\varpi_{zp}$  to lie close to zero showed a departure beyond  $1\sigma$  from the  $\mu_{RC}$  values found otherwise, indicating that  $\varpi_{zp}$  does have a measurable effect on the inferred RC luminosity. Finally, applying no strongly informative priors on the RC parameters nor  $\varpi_{zp}$  led to inferred values of  $\mu_{RC}$  and  $\varpi_{zp}$  being consistent with values in the literature, albeit with a large uncertainty of  $\sim 13 \mu\text{as}$  on the parallax zero-point offset, implying a fundamental limit on the uncertainty on this offset given the spatial parallax covariances.

Given that the choice of parallax zero-point offset did not dramatically affect the inferred luminosity of the clump (see Tables 9 and 10 and Fig. 8), we can reasonably use any value of  $\varpi_{zp}$  reported in the literature, including from this work, to attempt a calibration of seismology. Given the results for our runs on *Gaia* data with RC parameters constrained by seismology (Tables 6 and 8), we expect that  $\mu_{RC} = -1.634 \text{ mag}$  in *K* and in  $0.546 \text{ mag}$  in *G* would be roughly consistent with a temperature offset  $\Delta T_{\text{eff}}$  between  $\sim -100 \text{ K}$  and  $\sim -70 \text{ K}$  for temperatures in the APOKASC-2 catalogue (which, as has been noted, are already lower than those reported by Mathur et al. 2017, for the same stars). An offset of this size would fall within known systematic uncertainties on temperatures inferred from seismology (Slumstrup et al. 2019). However, it should be noted that this shift in temperature scale is degenerate with the scaling relations underestimating radii by  $\sim 2$  per cent compared to our estimates for radius using corrections by Sharma & Stello (2016), as discussed in Section 5.4.

In order to confirm these proposed shifts to temperature, we reran our asteroseismic model on our APOKASC-2 subsample for a range of temperature shifts extended down to  $-110 \text{ K}$  for both the *K* and *Gaia G* bands, with RC-corrected scaling relations. We found that when considering the *K* band, our calibration value for  $\mu_{RC}$  from *Gaia* corresponds to within  $1\sigma$  with a temperature shift of between  $-110$  and  $-70 \text{ K}$ . When considering the *G* band, the *Gaia*  $\mu_{RC}$  corresponds to within  $1\sigma$  for a shift between  $-70$  and  $-50 \text{ K}$ . Given that any calibrated correction to the temperature scale should be applied globally to the full APOKASC-2 subsample, we find that a temperature shift of  $-70 \text{ K}$  to the temperatures of our RC subsample of APOKASC-2 would produce seismic absolute magnitudes of the clump consistent with those found using *Gaia* DR2.

We only ran this test for the APOKASC-2 subsample, for which temperatures were all drawn from a uniform spectroscopic source. Since the temperatures for the full Y18 are not, claims about changes to temperature scales for this sample would be inappropriate.

The ability to make this inference reliably rests on our hierarchical treatment, as initially set out by H17, and treatment of the spatial correlations in parallax reported by Lindegren et al. (2018). As we improve our understanding of these correlations, our inferences using this and similar hierarchical models will improve. Similarly, it is known that population effects in age, metallicity, and temperature, among others, have an effect on the inferred luminosity of the RC Girardi (2016). Our hierarchical model can be further improved by accounting for these effects, as well as including parameters that check for consistent colours, as suggested by H17. As these hierarchical models improve in future work, so will our understanding of the RC, and our ability to calibrate asteroseismology.

## 6 CONCLUSIONS

Using two hierarchical models, based on the work by H17, we inferred the spread and position in absolute magnitude of a sample

of 5576 RC stars in the 2MASS *K* and *Gaia G* bands. We first did this using absolute magnitudes obtained through a completely distance-independent asteroseismic method, and probed systematics in asteroseismology by varying the temperatures of the sample, applying corrections to the scaling relations, and running our model on a subsample of stars with separate spectroscopic temperatures reported in APOKASC-2 (Pinsonneault et al. 2018). We then applied the results from seismology as strongly informative priors on the position and spread of the clump for our second hierarchical model. We applied this to *Gaia* DR2 data in order to see how the parallax zero-point varied, taking into account spatial correlations of parallaxes reported by Lindegren et al. (2018). We then applied strongly informative priors on the parallax zero-point in our *Gaia* model and allowed the RC parameters to roam more freely, to study the impact of published values for the zero-point offset on the RC. Finally, we performed a run of the *Gaia* model with no strongly informative priors on any parameters.

We leave the reader with the following conclusions:

- (i) By applying the H17 hierarchical model, with improvements to account for spatial correlations of parallaxes and marginalize over the parallax zero-point offset ( $\varpi_{zp}$ ), we find a mean value for  $\varpi_{zp}$  in the *Kepler* field to be  $-41 \pm 10 \mu\text{as}$  for our sample, consistent with all existing measures of  $\varpi_{zp}$  in the *Kepler* field. This offset results in an RC magnitude of  $-1.634 \pm 0.018$  in *K* and  $0.546 \pm 0.016$  in *G* for our sample.
- (ii) Applying a hierarchical model to our sample of absolute magnitudes obtained from asteroseismology, we find a spread of the RC in the 2MASS *K* band of  $\sim 0.03 \text{ mag}$  independent of our changes made to the sample, an order of magnitude lower than the value reported previously using *Gaia* TGAS parallaxes in H17. This extremely small spread highlights the power of seismology and the potential of the RC in the *K* band as a standard candle. In the *Gaia G* band we find a spread of  $\sim 0.13 \text{ mag}$  using APOKASC-2 temperatures, which is consistent with results found using *Gaia* DR2 parallaxes.
- (iii) We find that a small global change in temperature ( $\sim 10 - 20 \text{ K}$ ) can affect the inferred absolute magnitude of the RC from seismology by more than  $1\sigma$ , and is degenerate with the application of a correction  $f_{\Delta v}$  to the seismic scaling relations.
- (iv) We find values for the absolute magnitude of the RC from seismology to agree within  $1\sigma$  with those inferred from *Gaia* DR2 parallaxes in both the *K* and *G* bands, only if a global temperature shift of  $\sim -70 \text{ K}$  is applied to our RC subsample of APOKASC-2 stars. This shift is within expected systematic uncertainties on spectroscopic techniques. These differences are also degenerate with a shift in seismic radius of 2 per cent, which is within the uncertainty imposed by choice of corrections to the scaling relations.
- (v) A hierarchical Bayesian mixture model for a population of RC stars, as first set out by H17, continues to be an excellent tool for working with *Gaia* DR2 parallaxes, with the new additions of a parallax zero-point offset as a parameter and spatial correlations between parallaxes. Further additions will undoubtedly improve our inferences on RC stars, and with it, our ability to calibrate asteroseismology and *Gaia*.

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